

## PHOTOMETRIC AND RADIAL VELOCITY VARIATIONS OF RCrB NEAR MAXIMUM LIGHT

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**ABSTRACT:** Fourier analysis of the light curves of RCrB in V band near maximum shows that in addition to several significant short periods there is a modulation of the visual light with a period around 1170 day, similar to that of L band flux, noticed by Strecker. This indicates that there is some contribution to the visual light variations of the star from the pulsating circumstellar dust. Radial velocities of R CrB obtained at Kavalur during February-May 1985 show variations with a period around 47 days.

### 1. INTRODUCTION

In the case of R CrB, the light near maximum fluctuates with an amplitude in the range of about 0.1-0.3 mag and appears very irregular in shape, amplitude and period. In order to check the possibility that the apparent irregular behaviour is as a result of the superposition of several independent waves, we decided to Fourier analyse the available light curves of R CrB obtained near maximum. The technique of Discrete Fourier Transform as developed by Deeming (1975) was employed, since the information on aliasing due to the different data spacings can be obtained directly from the spectral window derived from the times of observations. All the available measurements in V band from 1971 to 1983, except those belonging to either a primary or a secondary fading, were included in the analysis.

### 2. POWER SPECTRUM

Several peaks of significant and comparable amplitudes could be identified in the power spectrum. We made a least square solution for the determination of the optimum periods and their corresponding amplitudes and phases of the dominant peaks. We found that with these quantities the observed light curves could not be reproduced satisfactorily. This prompted us to suspect that the various periodicities had their origins from different data segments obtained at different epochs.

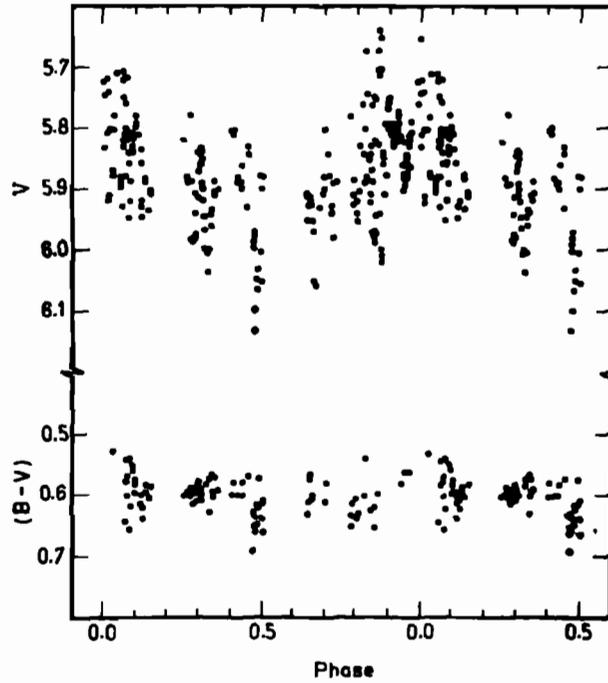


Fig. 1a. V and B-V values of R C rB near maximum during 1971-83 folded over a 1170 day period.

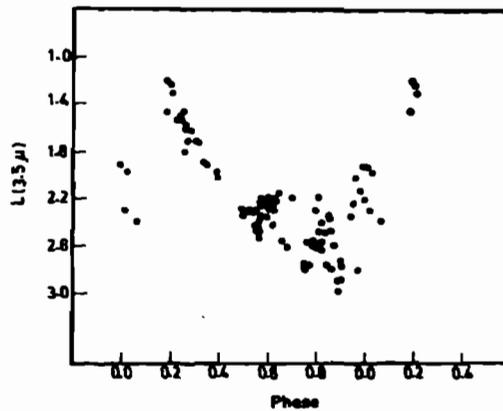


Fig. 1b. L magnitudes available during 1968-82 folded over a 1170 day period.

Hence, we made a piece-wise analysis of the data which confirmed our suspicion. Further, we could find no single period which is present in all the different data segments. These results convincingly rule out the possibility that the observed irregular light fluctuations are caused by the superposition of several independent oscillations.

### 3. LONG-TERM BEHAVIOUR

#### 3.1 Visual light

Of the several periods present in the power spectrum, the most dominant is a long-term modulation with a period around 1200 day. In Fig. 1a, we have plotted the V and B-V magnitudes of R CrB obtained near maximum against the photometric phases calculated using the following ephemeris:

$$JD = 2444294.0 + 1170.0 E$$

The initial epoch corresponds to the time of maximum brightness in V for the 1170 day modulation. Both the period and the initial epoch are derived using least square method assuming a sinusoidal form for the variation. The total amplitude in V is found to be around 0.15 mag. Apparently, (B-V) does not show significant variation over this period. It is clear from Fig.1a, that the short-term fluctuations are superposed on this long-term behaviour.

#### 3.2 L Band (3.5 $\mu$ ) flux

The evidence of long-term modulation of the light draws one's attention towards the variability of infrared excess at L band in R CrB, noticed first by Humphreys and Ney (1974), and confirmed later by Strecker (1975) who derived a characteristic time of about 1100 days for the variation. The available L band observations from 1968 to 1982 are plotted in Fig.1b with the above ephemeris. It is evident from the figure that the period fits the infrared data also fairly well. The variability in infrared excess is attributed to the pulsating circumstellar dust around R CrB (Feast et al. 1977). If true, this implies that there is some contribution to the optical variation from the circumstellar dust.

### 4. RECENT RADIAL VELOCITY DATA

In order to look for the suspected pulsation, R CrB was observed spectroscopically on 27 nights during February-May 1985 with the 75-cm Telescope at Kavalur at a dispersion of 26  $\text{\AA} \text{ mm}^{-1}$ . In Fig.2a, we have plotted the radial velocities derived by us against the corresponding Julian Days of observation. It is clear from the figure that the radial velocity of the star is variable, but no definite pattern for the variation could be seen. The total amplitude of variation is around 8  $\text{kms}^{-1}$  in about 50 days. We have obtained

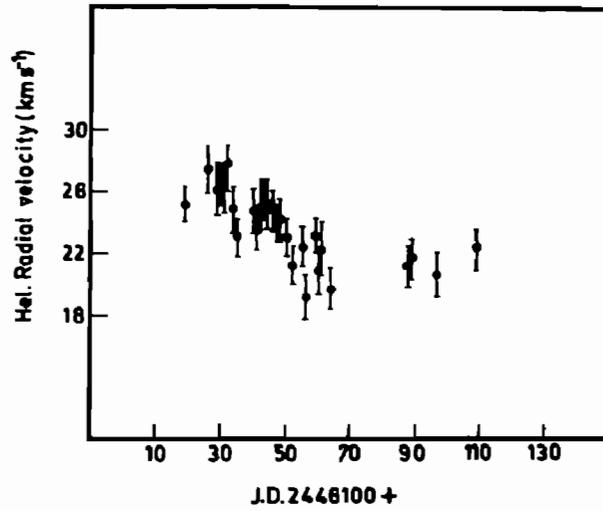


Fig. 2a. Radial velocities of R CrB obtained at Kavalur during February-March 1985.

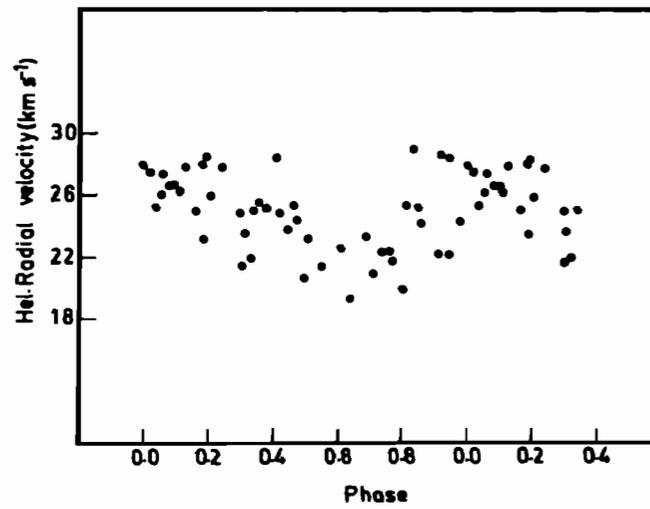


Fig. 2b. Radial velocities of R CrB near maximum available since 1972 folded over a 47.18 day period.

a few radial velocity measurements during 1983 also. A period of 47.18 day is found to fit fairly well all our observations along with those given by Fernie et al. (1972), as shown in Fig.2b. However, the period is a poor presentation if observations earlier to 1972 are also included. An extended version of this paper will be published elsewhere.

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