

A Low Frequency Radio Telescope at Mauritius for a Southern Sky Survey

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Abstract. A new, meter-wave radio telescope has been built in the north-east of Mauritius, an island in the Indian Ocean, at a latitude of -20.14° . The Mauritius Radio Telescope (MRT) is a Fourier Synthesis T-shaped array, consisting of a 2048 m long East-West arm and an 880 m long South arm. In the East-West arm 1024 fixed helices are arranged in 32 groups and in the South arm 16 trolleys, with four helices on each, which move on a rail are used. A 512-channel digital complex correlation receiver is used to measure the visibility function. At least 60 days of observing are required for obtaining the visibilities up to 880 m spacing. The Fourier transform of the calibrated visibilities produces a map of the area of the sky under observation with a synthesized beam width $4' \times 4.6' \sec(\delta + 20.14^\circ)$ at 151.5 MHz.

The primary objective of the telescope is to produce a sky survey in the declination range -70° to -10° with a point source sensitivity of about 200 mJy (3σ level). This will be the southern sky equivalent of the Cambridge 6C survey. In this paper we describe the telescope, discuss the array design and the calibration techniques used, and present a map made using the telescope.

Key words. Radio telescope—low frequency—imaging—southern sky—Fourier Synthesis.

1. Introduction

Surveying the sky and compiling catalogs of celestial objects has been a major part of astronomical research for centuries. The first systematic survey of the radio universe was carried out by Grote Reber (1944) using a backyard telescope with a resolution of 12° operating at a frequency of 160 MHz. With the quest for higher angular resolution the exercise of surveying soon shifted to higher frequencies. Even so, many low frequency surveys were carried out after Reber's survey some of which are summarized in Table 1. The table clearly indicates that the sixth Cambridge survey (6C) (Baldwin *et al* 1995) is by far the most extensive

Table 1. Surveys below 1 GHz.

Northern declination surveys					
Freq.	Observatory	Resol.	Dec. coverage	Sensit.	No. of sources
408 MHz	Effelsberg	$0.85^\circ \times 0.85^\circ$	$-10^\circ + 50^\circ$	0.2 Jy	—
178 MHz	Cambr. 3CR	$2' \times 2'$	$-5^\circ + 90^\circ$	9 Jy	328
178 MHz	Cambr. 4C	$7.5' \times 7.5'$	$-7^\circ + 80^\circ$	2 Jy	4843
38 MHz	Cambr. WKB	$45' \times 45'$	$-45^\circ + 35^\circ$	14 Jy	1000
34.5 MHz	GEE TEE	$30' \times 30'$	$-30^\circ + 60^\circ$	5 Jy	≈ 3000
151 MHz	Cambr. 6C	$4' \times 4'$	$+30^\circ + 90^\circ$	0.12 Jy	$> 10^5$
Southern declination surveys					
38 MHz	Cambr. WKB	$45' \times 45'$	$-45^\circ + 35^\circ$	14 Jy	1000
34.5 MHz	GEE TEE	$30' \times 30'$	$-30^\circ + 60^\circ$	5 Jy	≈ 3000
408 MHz	Parkes	$0.85^\circ \times 0.85^\circ$	$-60^\circ + 10^\circ$	1 Jy	—
408 MHz	Molongolo	$2' \times 2'$	$-60^\circ + 18^\circ$	0.6 Jy	> 12000
843 MHz	MOST	$0.4' \times 0.4'$	$-90^\circ - 30^\circ$	—	ongoing
151.5 MHz	MRT	$4' \times 4'$	$-70^\circ - 10^\circ$	0.2 Jy	ongoing

survey at low frequencies. This survey provides a moderately deep radio catalog reaching a source density of about $2 \times 10^4 \text{ sr}^{-1}$ over most of the sky north of $\delta = +30^\circ$ with an angular resolution of $4.2' \times 4.2' \text{ cosec}(\delta)$ and a limiting flux density of 120 mJy at 151 MHz. An equivalent of the 6C survey for the southern sky does not exist.

Since the survey of Mills *et al.* (1958) at 80 MHz there has not been much effort at low frequencies to map the southern sky apart from the Parkes 408 MHz survey. The Culgoora (Slee 1977) observations at 80 and 160 MHz were made mainly to study known sources and determine their spectral indices. The largest existing radio survey of the southern sky is the Parkes-MIT-NRAO survey (Griffith & Wright 1993) at 5 GHz. There is an obvious need to survey the southern sky at a frequency around 150 MHz. At this frequency synchrotron sources show up much better than at higher frequencies, such as 408 MHz, due to their spectra. Sources also show up better at 150 MHz than at lower frequencies since the absorption due to the interstellar gas is much less than at deca-meter wavelengths. For this purpose a radio telescope operating at 150 MHz has been constructed at Bras d'Eau, Mauritius.

The Mauritius Radio-Telescope (MRT) has been constructed and is operated collaboratively by the Raman Research Institute, the Indian Institute of Astrophysics and the University of Mauritius. It is situated in the North-East of Mauritius (latitude 20.14° South, longitude 57.74° East), an island in the Indian Ocean. It is a T-shaped array with an East-West (E-W) arm of length 2048 m having 1024 helical antennas and a South (S) arm of length 880 m consisting of a rail line on which 16 movable trolleys each with four helical antennas are placed (Fig. 1). The helical antennas respond to frequencies between 80 and 160 MHz. Presently the telescope is operated at 151.5 MHz which allows maximum interference-free observations. The helices are mounted with a tilt of 20° towards the south so that they point towards a declination of -40° at the meridian. The declination coverage corresponding to the Half Power Beam Width (HPBW) of the helices is from -70° to -10° . Fig. 2 shows an aerial view of the telescope and Fig. 3 shows a closer view of the helices used.

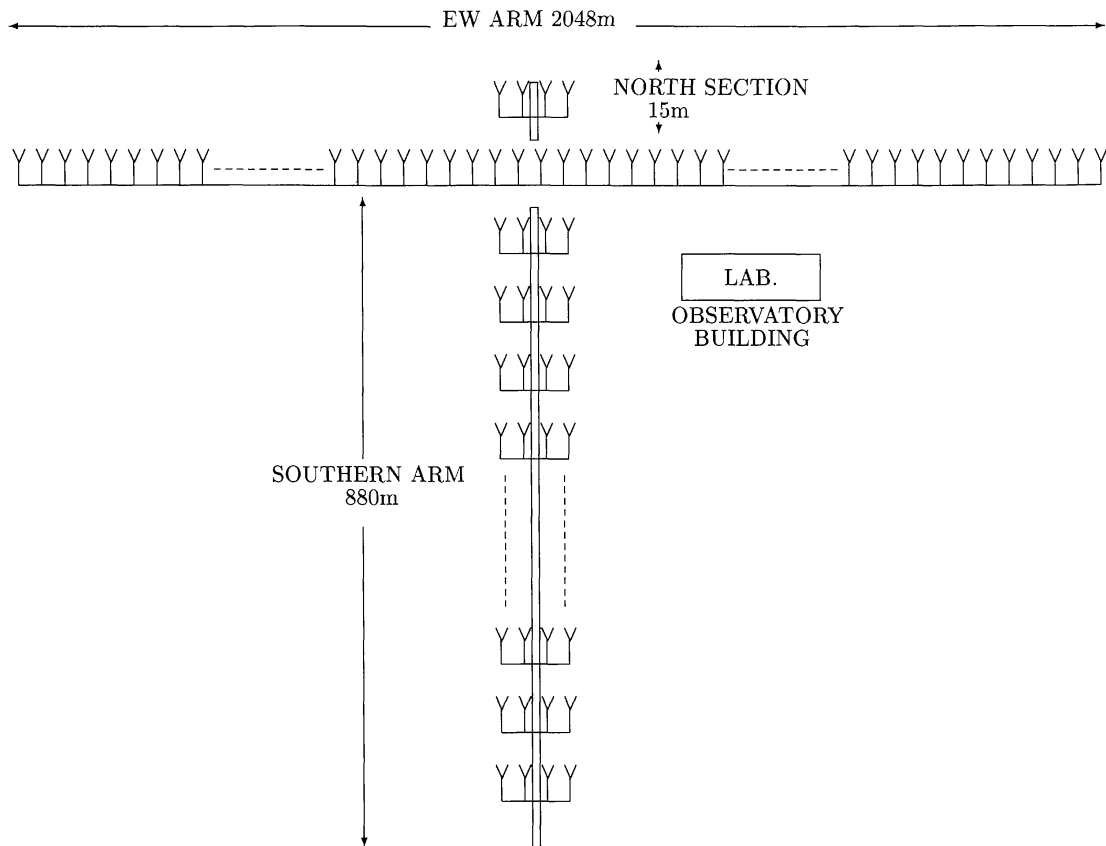


Figure 1. A schematic of the array. The E-W arm has 1024 helices which are divided in 32 groups of 32 helices each. The N-S arm has 16 trolleys each with 4 helices.

The E-W arm is divided into 32 groups of 32 helices each. All the E-W groups are not at the same height, a situation imposed by the uneven terrain. Each trolley in the S arm constitutes one S group. Both E-W and S group outputs are heterodyned to an intermediate frequency (IF) of 30 MHz, using a local oscillator (LO) at 121.6 MHz. The 48 group outputs are then amplified and brought separately to the observatory building via coaxial cables. In the observatory, the 48 group outputs are further amplified and down-converted to a second IF of 10.1 MHz. The 32 E-W and 16 S group outputs are fed into a 32×16 complex, 2-bit 3-level digital correlator sampling at 12 MHz. The 512 complex visibilities are integrated and recorded at intervals of 1 second. At the end of 24 hours of observation the trolleys are moved to a different position and new visibilities are recorded. A minimum of 60 days of observing are needed to obtain the visibilities up to the 880 m spacing.

The Fourier transform of the phase corrected visibilities obtained after the complete observing schedule produces a map of the area of the sky under observation with a synthesized beam width of $4' \times 4.6' \sec(\delta + 20.14^\circ)$. The phase corrections mentioned above take into consideration the non-coplanarity of the baselines. The expected root mean squared (RMS) values of the background in the synthesized images, arising from the system noise with a 1 MHz bandwidth and an integration time of 8 seconds and from the confusion noise are expected to be around 200 mJy (3σ) and 10 mJy respectively.

Table 2 summarizes the MRT specifications.

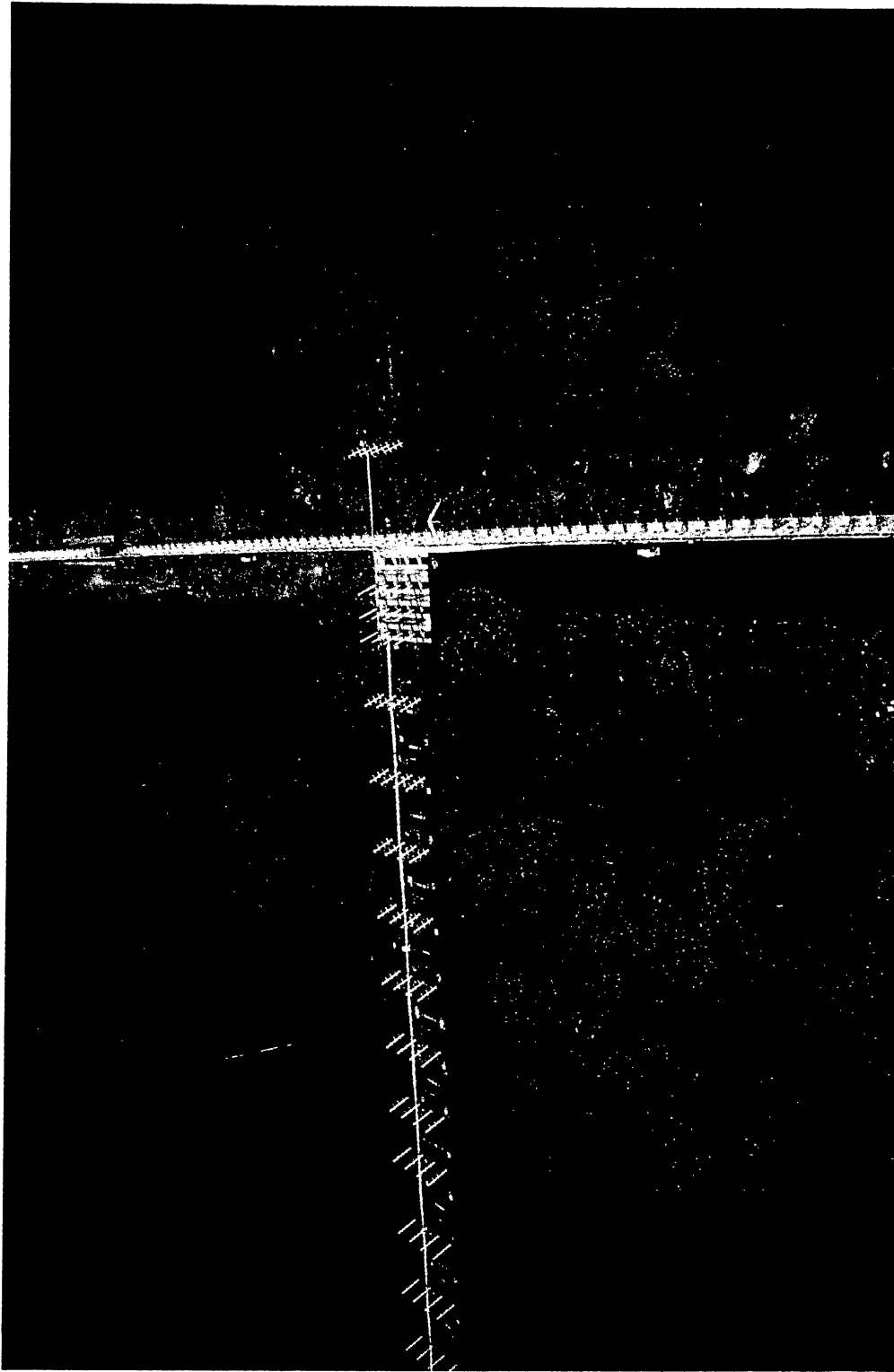


Figure 2. An aerial view of the Mauritius Radio Telescope.

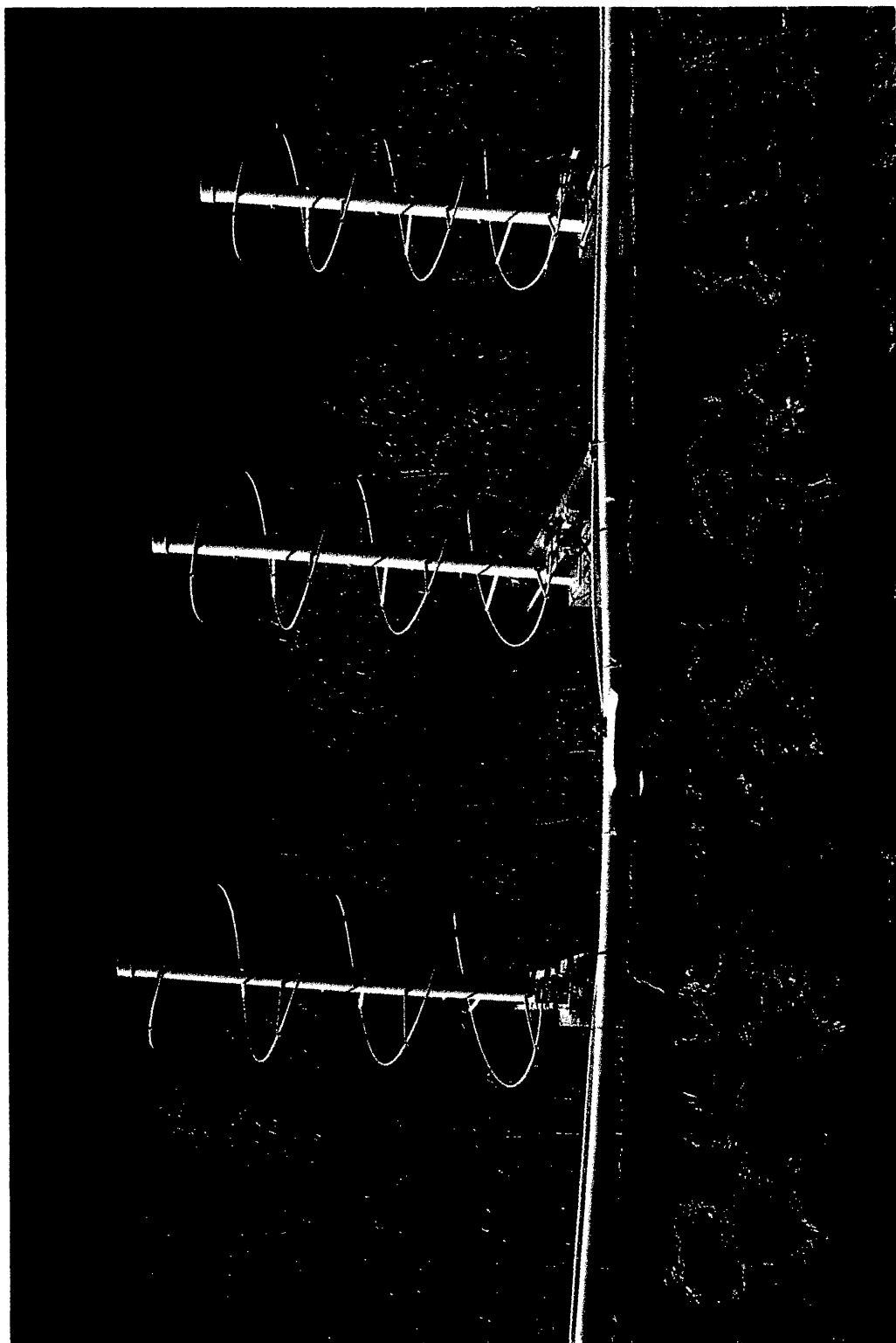


Figure 3. A closer view of the helices used.

Table 2. MRT specifications.

Observing frequency	151.5 MHz.
Telescope configuration	T-shaped 2048 m E-W arm and 880 m N-S arm.
Basic element	Helical antenna.
Polarization	Right Circular.
HPBW of helix	$60^\circ \times 60^\circ$.
Declination coverage	-70° to -10° .
Collecting area of helix	4 m^2 at 150 MHz.
East-West arm	32 groups with 32 helices each.
North-South arm	15 trolleys each with 4 helices.
1st IF frequency	30 MHz.
2nd IF frequency	10.1 MHz.
Instrumental bandwidths	0.15, 1.0, 1.5, 3.0 MHz.
Digitization before correlation	2-bit 3-level.
Correlation receiver	32×16 complex channels.
No. of baselines measured per day	32×16 .
Minimum and maximum baselines	0, 1024λ .
Time to get full resolution image	60 days.
Synthesized beam-width	$4' \times 4.6' \text{ sec}(\delta + 20.14^\circ)$.
Point source sensitivity	200 mJy (3σ).

2. The telescope

2.1 Design criteria

The array was designed by considering the availability of the 512 channel Clark Lake correlator system¹ (Erickson *et al* 1982), the constraints due to the available terrain and the presence of man-made interference.

Configuration: A T-shaped configuration was chosen for its simplicity. Aperture synthesis with fixed antennas in the E-W arm and movable elements (trolleys) in the S arm were chosen to minimize the hardware required. An abandoned old railway line running North-South was rebuilt for use as the South arm. This new rail line, slopes downwards at about $1/2^\circ$ to the horizontal till about 655 m, and then slopes upwards at about 1° to the horizontal. On this rail the trolleys cannot approach the E-W arm nearer than 11 m.

To ensure that the array responds to structures of all sizes in the sky, the array should provide all spacings available in a square aperture. To meet this requirement at MRT, a 15 m North extension with one trolley almost touching the E-W arm has been built and is used to measure the low spatial frequencies. However, this trolley cannot approach the E-W array nearer than 2 m. Hence baselines with 1 m spacing in the S direction are not measured. Non-zero E-W baselines with zero-spacing along the S direction are obtained by multiplying the groups of the eastern arm with the output of four helices which are a part of the first group (closest to the center) of the western arm. This has the same primary beam as a trolley and ensures a weighting similar to baselines with non-zero spacing along the S direction.

¹ After the closure of the Clark Lake radio-telescope the two-bit three-level 512 channel complex correlation receiver was kindly donated to the MRT.

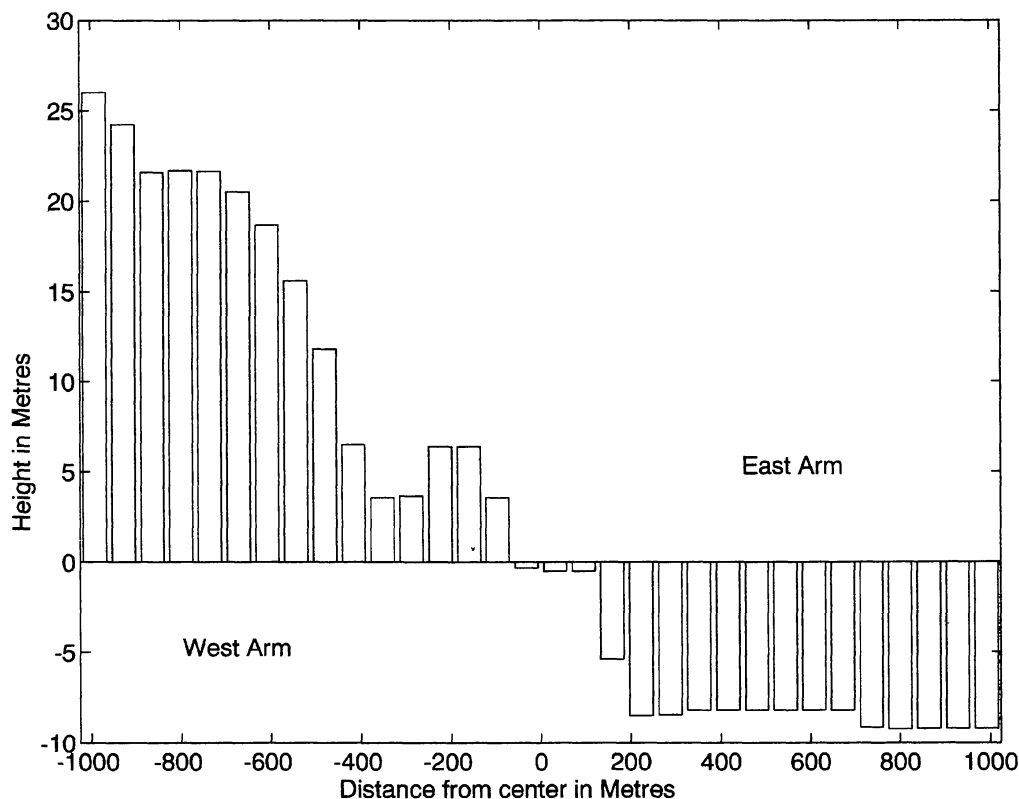


Figure 4. The E-W height profile. The maximum height difference in the E-W arm is 35 m.

Terrain: The local terrain is rocky and very uneven, especially along the E-W arm with height differences of up to 35 m. To minimize the problems of non-coplanarity, it was decided to level the E-W arm in multiples of 64 m (32λ) so that the antennas in each group will be at the same height. The height profile of the groups is shown in Fig. 4. For optimal use of the 512-channel complex correlation receiver, several schemes were considered before arriving at a configuration consisting of 16 movable trolleys in the N-S arm and 32 groups in the E-W arm. Details of these considerations are given by K Golap (1998).

Interference: When this telescope was conceived, man-made interference in Mauritius was very low. By the time the array became operational, the increased use of many local communication networks, had increased man-made interference. The front end of the receiver system has been built with sufficient bandwidth so that the observing frequency can be shifted (within 145–155 MHz) to an interference-free zone by tuning the LO. The 1 MHz band around 151.5 MHz has been found to be relatively quiet and is therefore presently being used.

2.2 The MRT array

Helix: The primary element is a peripherally fed monofilar² axial-mode helix of 3 turns with a diameter of 0.75 m and a height of 1.75 m (Fig. 5). This is mounted above

² Monofilar: A term used to distinguish single conductor helix from helices with two or more conductors.

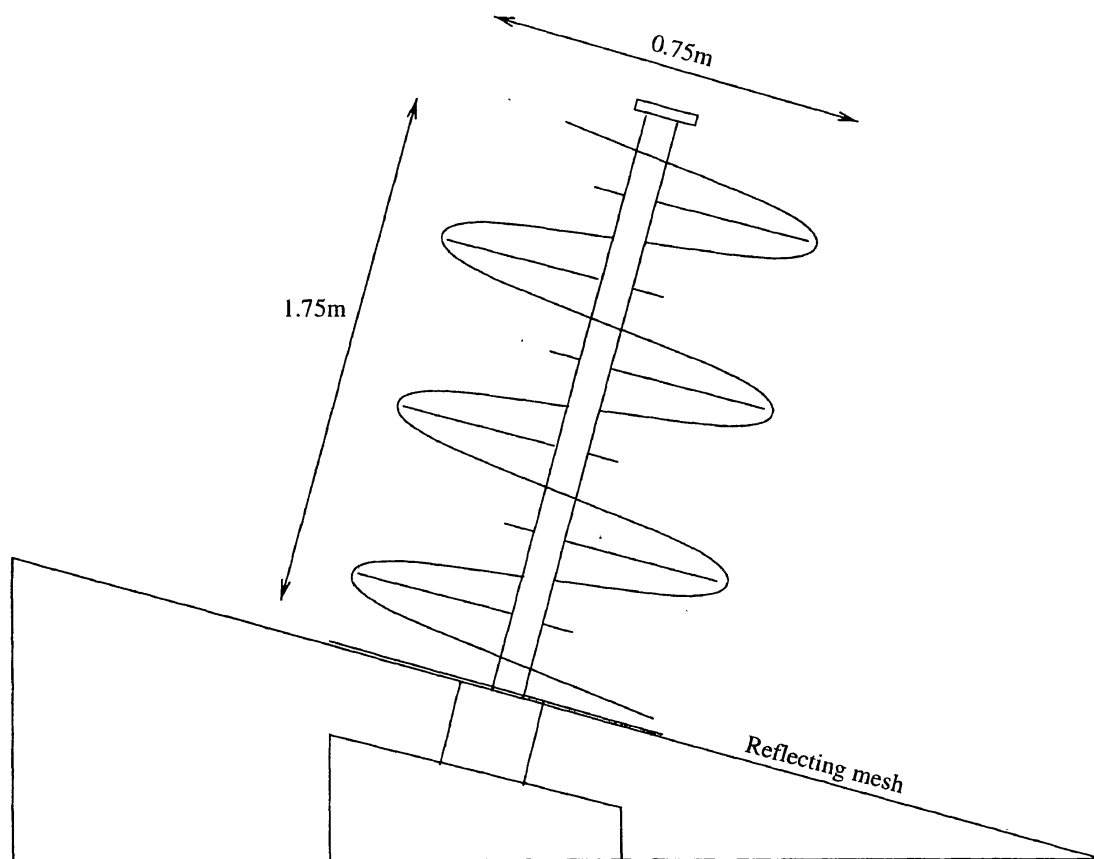


Figure 5. The MRT helix is a peripherally fed monofilar axial mode helix of 3 turns. The helices are mounted with a tilt of 20° to allow a better coverage of the southern sky.

a stainless steel reflector mesh of grid size $2'' \times 2''$. The helix is wound using a 1.5 cm diameter round aluminium tubing supported with a central UV (Ultra Violet) stabilized, light-weight PVC (Poly Vinyl Chloride) cylinder and radial PVC rods. The axial mode provides maximum radiation along the helix axis.

The helix responds to frequencies between 100 and 200 MHz with right circular polarization (IEEE definition). A quarter-wave transformer used in the feed network optimizes the VSWR to ≤ 1.5 around 150 MHz. The helical antenna, with its reflector, has a canonical collecting area of about λ^2 (4 m^2 at 150 MHz) with a HPBW of about $60^\circ \times 60^\circ$. The helices are mounted with a tilt of 20° towards the South to get a better coverage of the southern sky (-70° to -10° dec) including the southern-most part of the galactic plane, a region largely unexplored at meter wavelengths.

E-W and S arms: The E-W arm consists of 1024 helices mounted on a 2 m wide ground plane with an inter-element spacing of 2 m (λ at 150 MHz) and is divided into 32 groups of 32 helices each. As already mentioned, due to the uneven terrain all the E-W groups are not at the same height. The HPBW of the primary beam of each E-W group is $1.8^\circ \times 60^\circ$ and allows observation of a source for roughly $7 \times \sec(\delta)$ minutes. In each group, four helices are combined using power combiners and the combined output is pre-amplified in a low noise amplifier (noise temperature $\approx 300\text{K}$). Eight such amplified outputs are further combined and amplified to form a group output in the E-W array.

Four helices mounted on a trolley with a 4 m-wide ground plane constitute a group in the S array with a primary beam of HPBW $15^\circ \times 60^\circ$. Each of the 32 E-W and 16 S group outputs are heterodyned to an IF of 30 MHz, amplified and transmitted, using coaxial cables, to the observatory building situated close to the center of the array. Equal lengths of coaxial cables are used irrespective of the distance of the groups from the observatory to ensure that the interferometer outputs are not affected much by changes in the ambient temperature. The configuration of the E-W group is shown in Fig. 6. Each N-S trolley has a similar signal flow path except that there are only 4 helices and 1 pre-filter unit.

Local Oscillator System: Low loss (2.4 dB/100 m at 120 MHz) heliax cables running parallel to the East, West and South arms of the array are used to distribute the LO signal required for heterodyning. The 121.6 MHz LO signal at each group is obtained by using directional couplers. To maintain a more or less constant power level for the LO signal (-5 dBm to -7 dBm) across the array, broad-band high-power amplifiers are placed at several places along the heliax cable.

To minimize the spurious correlations in the visibility measurements, the following hardware features are built into the LO system.

- The first LO is phase switched to minimize the effects of cross coupling. The Clark Lake correlator has the facility for generating the Walsh functions required for switching. But it is not practical to use many independent switching waveforms in a series fed LO system owing to the problem of distribution. Hence, we use a simple arrangement in which the LO to the first eight southern groups remain unswitched while, the eastern, the western and the remaining southern groups are switched with orthogonal square waves. The disadvantage of this scheme is that the products of the outputs from groups having the same switching signal are prone to spurious correlations (e.g., the product of any two eastern arm group outputs).
- The signal generator used to produce the LO also generates spurious signals at other frequencies. The spurious signals around 30 MHz, which is our first IF frequency, leaks to the IF port from the LO port of the mixer and results in correlation several times the detection limit. This is minimized by using a band rejection device centered around 30 MHz in the LO path. Band rejection is achieved using a band-pass filter centered around 30 MHz and a power-splitter (Fig. 6). The filter absorbs the signal around 30 MHz and reflects the LO required for heterodyning.

2.3 The receiver system

2.3.1 Second IF and correlator modules

In the observatory building the 48 group outputs are further amplified. Each of the 32 E-W group outputs are split into two using power dividers. One set of the outputs is combined to form a fan beam of $4' \times 60^\circ$ using phase-shifter modules. This single beam forming receiver provides a two degree tracking system. Details of pulsar observations carried out using this system are described by N. Issur (1997).

The second set of 32 E-W group outputs and the 16 S outputs are down-converted to 10.1 MHz (second IF). Four IF bandwidths, ranging from 0.15 to 3 MHz are selectable. Signals are then fed to an Automatic Gain Control (AGC) module which keeps the output power level constant. Each of the 16 group outputs of the S arm and

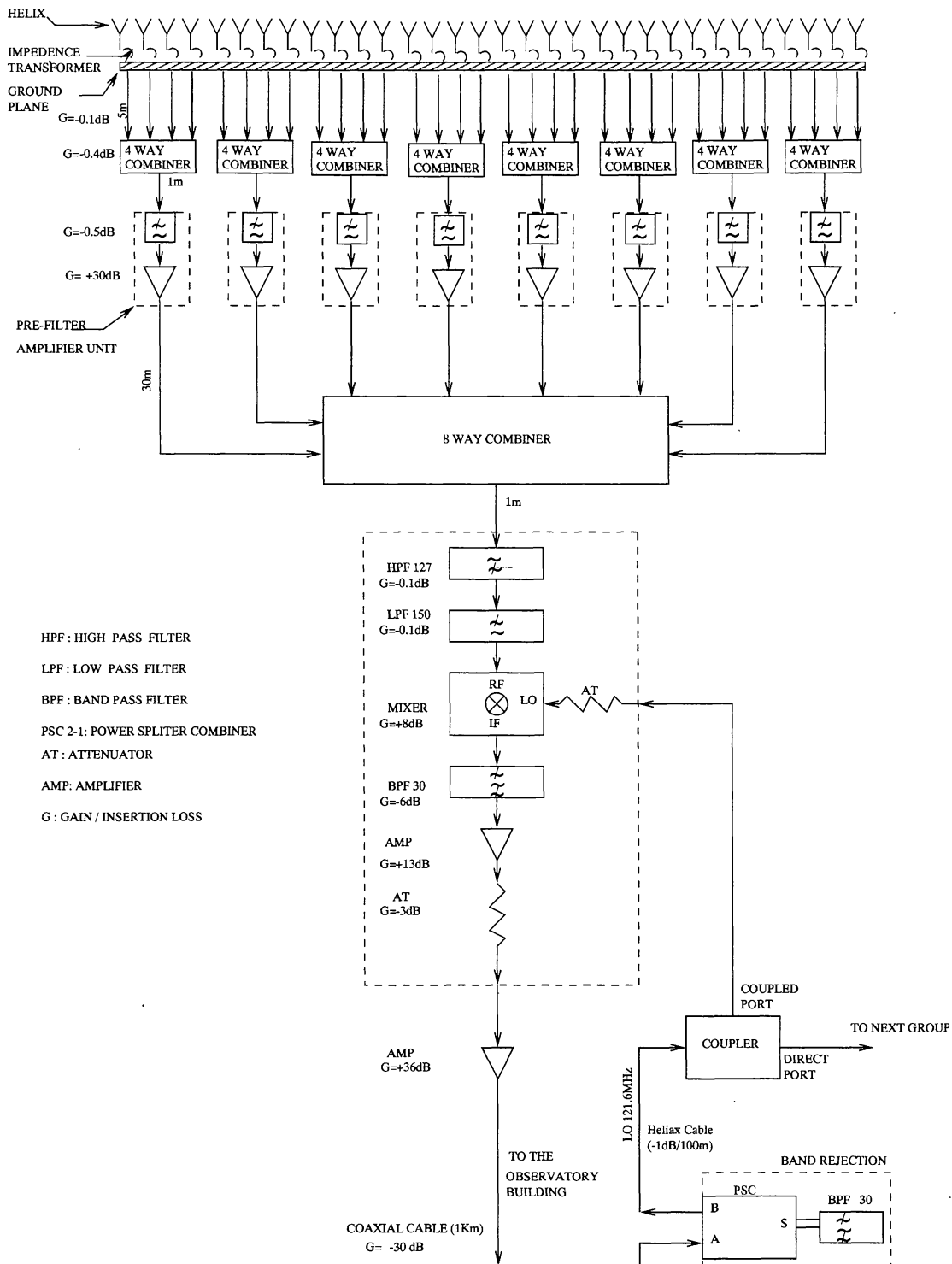


Figure 6. Signal path of an E-W group. The N-S trolley has a similar flow path except that there are only 4 helices and 1 pre-filter unit.

the 16 group outputs of the E arm is split in a quadrature hybrid to obtain the in-phase and quadrature components. Then the signals are digitized to 2-bit 3-level, sampled at 12 MHz and fed into the 32×16 complex correlator to obtain the E-W \times S outputs. The high sampling rate is to reduce the loss of sensitivity due to quantization. More

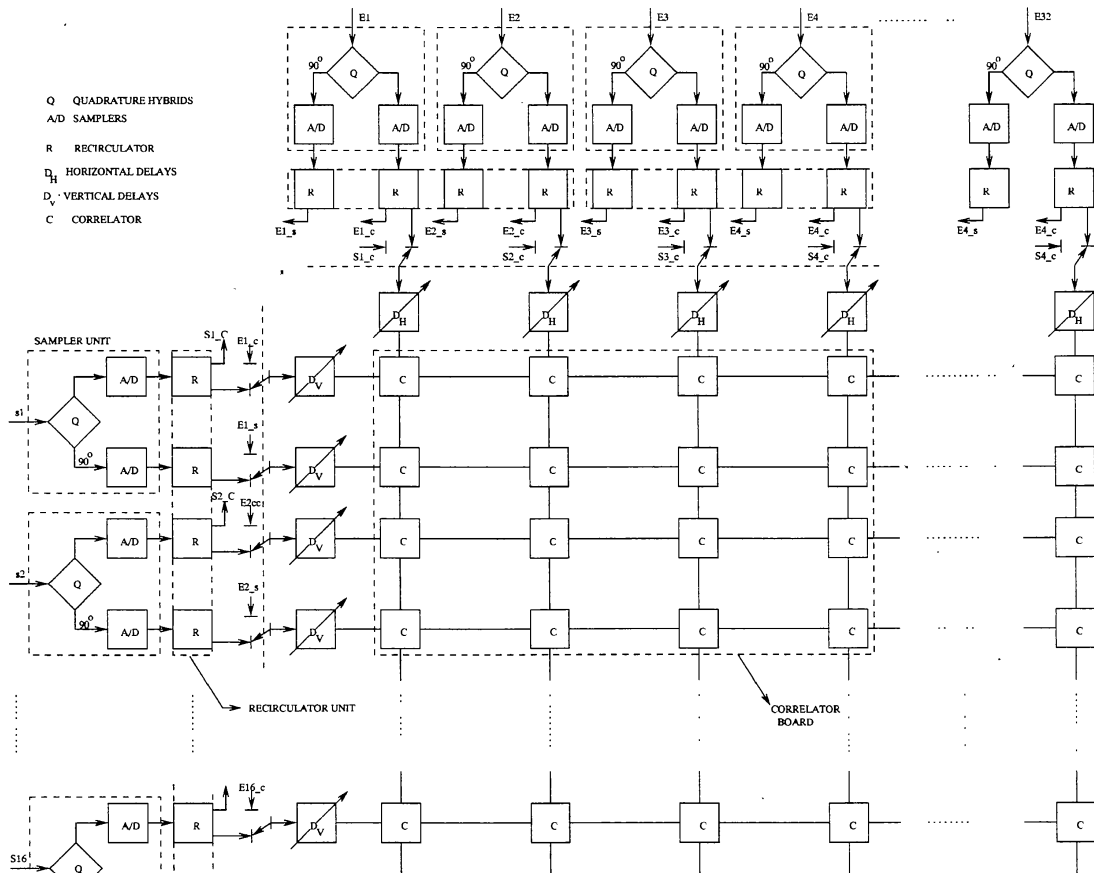


Figure 7. Diagram showing the general layout of the signals to the correlators.

details of the system are given by Erickson *et al.* (1982). Fig. 7 shows the general layout of the samplers, delay lines and correlators.

The last group of the East array (E16) is fed to the correlator in the place of the 16th trolley. This gives a set of baselines formed between E16 and the E-W array on all observing days. This helps to check the repeatability of data. These baselines with each trolley give 31 independent closure information which are used in the calibration. This mode of observing reduces the number of usable trolleys to only 15.

Many observing programs require the formation of fan beams (Issur 1997) corresponding to each arm of the array. The correlator can also be configured to obtain E \times E-W and S \times S correlations.

Self-correlators: Sixty four self-correlators are available in the receiver system. This provides the necessary information for obtaining the normalized correlation coefficient from the measured digital correlation counts using the correction given by D'Addario *et al.* (1984).

$$\rho = \gamma - \frac{\gamma^3(\alpha_1^2 - 1)(\alpha_2^2 - 1)}{6} \quad (1)$$

where $\gamma = (\pi/2)(N/N_{\max}) \exp(\alpha_1^2 + \alpha_2^2)/2$. α_1 and α_2 are the V_{th}/σ of the two channels being correlated and N/N_{\max} is the ratio of the correlation value to the maximum value possible out of the 2-bit 3-level system (V_{th} is the threshold voltage for digitization and σ is RMS value of the signal fed to the digitizer).

The AGCs maintain a constant power level to the samplers, even though the brightness distribution of the sky changes. Therefore we do not get the amplitude information of the signal from the sky. This results in identical correlation for a weak source in a weak background and a stronger source in a correspondingly stronger background. To measure the absolute brightness while using the AGC there is a need to measure the gains. In many telescopes this is done by adding a fixed amount of a known signal at each antenna. At the MRT, however, the variation in the background radiation as seen by the E-W and the S groups are measured separately by switching off the AGCs, one in each of the E-W and the S arms and using the self-correlators to measure the total power output of these groups.

The self correlators of the MRT are wired such that they measure the probability (P) that the input signal amplitude, V , is between the threshold levels used for digitization. This probability for a zero-mean Gaussian signal with RMS fluctuation of σ and a symmetric 2-level digitizer with voltage threshold levels $\pm V_{th}$, is given by

$$P = \frac{1}{\sigma\sqrt{2\pi}} \int_{-V_{th}}^{+V_{th}} e^{-(V/\sqrt{2}\sigma)^2} dV = \text{erf}\left(\frac{V_{th}}{\sqrt{2}\sigma}\right), \quad (2)$$

knowing P , the RMS fluctuation σ of the signal can be obtained. The analog correlation ρ_a is obtained using the relation

$$\rho_a = \rho \times \sigma_1 \times \sigma_2 \quad (3)$$

where σ_1 and σ_2 are the RMS of the signals correlated.

In a 2-bit 3-level correlator, sampling the digitized signal at Nyquist rate, the maximum sensitivity obtainable relative to an analog correlator is 0.81. This is obtained when V_{th}/σ is 0.61. The sensitivity changes by only 5% from this optimal value when the signal power changes by 40% (Bowers & Klinger 1974). Hence, switching off the AGC does not affect the signal-to-noise ratio (SNR) of the channels used to measure the total power as the variation of the sky brightness is less than 40%.

Recirculator: Although the use of larger bandwidths results in better sensitivity of a telescope, it restricts the angular range over which an image can be made if the relative delays between the signals being correlated are not compensated. When the uncompensated delay between the signals becomes comparable to the inverse of the bandwidth used, the signals will be decorrelated. At MRT we normally use a bandwidth of 1 MHz. Since the E-W group has a narrow primary beam of two degrees in RA, this bandwidth does not pose a problem for synthesizing the primary beam in this direction. However both the E-W and N-S groups have wide primary beams in declination extending from -70° to -10° . For signals arriving from zenith angles greater than 10 degrees on N-S baselines longer than 175 m, the uncompensated delay results in bandwidth decorrelation greater than 20%.

To overcome this loss of signal one has to measure visibilities with appropriate delay settings, while imaging different declination regions. To keep the loss of signal for a bandwidth of 1 MHz to less than 15% in the entire declination range -70° to -10° , the longer baselines have to be measured with four delay settings.

With the existing correlator system this implied observations for four days at trolley locations measuring long baselines. Furthermore, to be able to make interference free maps, observations have to be repeated approximately 3 times at each location. This would make the time required to complete the survey very large.

To solve this problem, a recirculator system which measures visibilities with different delay settings using the available correlators in a time multiplexed mode has been incorporated in the receiver system (Sachdev 1998). In this system, the data are sampled at 3 MHz, stored in a buffer memory and the correlations are measured at 12 MHz rate. Such processing by the correlator at a higher speed than the input sampling rate allows correlations to be measured with four delay settings. This ensures that the observations at each trolley allocation, even at longer baselines, can be carried out in one day, thereby, improving the surveying sensitivity. The loss of sensitivity due to the decrease in the input sampling rate from 12 MHz to 3 MHz is not serious. The recirculator system is described in detail by Sachdev (1998).

The recirculator system can configure the correlator in the $(E + W) \times E$ and $S \times S$ modes for one integration time (roughly 100 ms in every second of data collection). Since both E-W and S arrays are linear, this mode provides visibilities on several redundant baselines to calibrate the array using the Redundant Baseline Calibration technique (Noordam & Bruyn 1982). Presently the software at MRT does not make use of this data for array calibration.

2.4 Observations

In the T configuration, we are interested in sampling baselines with North-South components from 0 m to 880 m with a sampling of 1 m (except the 1 m spacing [section 2.1]). This ensures that the grating response will be confined well outside the primary beam response. To measure the visibilities with a 1 m spacing using 15 trolleys however requires 60 days of observing. Each day the trolleys are spread over 84 m with an inter-trolley spacing of 6 m. 512 complex visibilities are recorded with 1 sec of integration. After obtaining data for 24 sidereal-hours the trolleys are moved by one meter. After 6 days of successful observing we get visibilities measured over 90 m. The trolleys are then moved *en bloc* to measure the next set of visibilities.

3. Imaging with the MRT

The visibilities measured are processed off-line using MARMOSAT³, the MAuRitius Minimum Operating System for Array Telescopes (Dodson 1997). This is designed in-house to transfer the visibilities to images which can be ported to AIPS. The flow diagram showing the various processing stages is shown in Fig. 8.

The programs in MARMOSAT cater to the following needs: detection of bad data, determination of complex gain of the antennae, position calibration of antennas, combining data of different days, transforming the visibilities to get brightness distribution.

The program **Online** presents an uncalibrated view of the data. It is designed to read visibilities from several baselines and offers to display many functions such as phase, magnitude, closure and fringe frequencies of the visibility measured. The program **Health** monitors the signal-to-noise ratio of the visibilities, correlator offsets, magnitudes of the visibility of all the baselines and displays the mean and the scatter. This helps to detect gross errors quickly. The program **Interference Detection** detects the

³ A marmosat is a small monkey, distantly related to the ape.

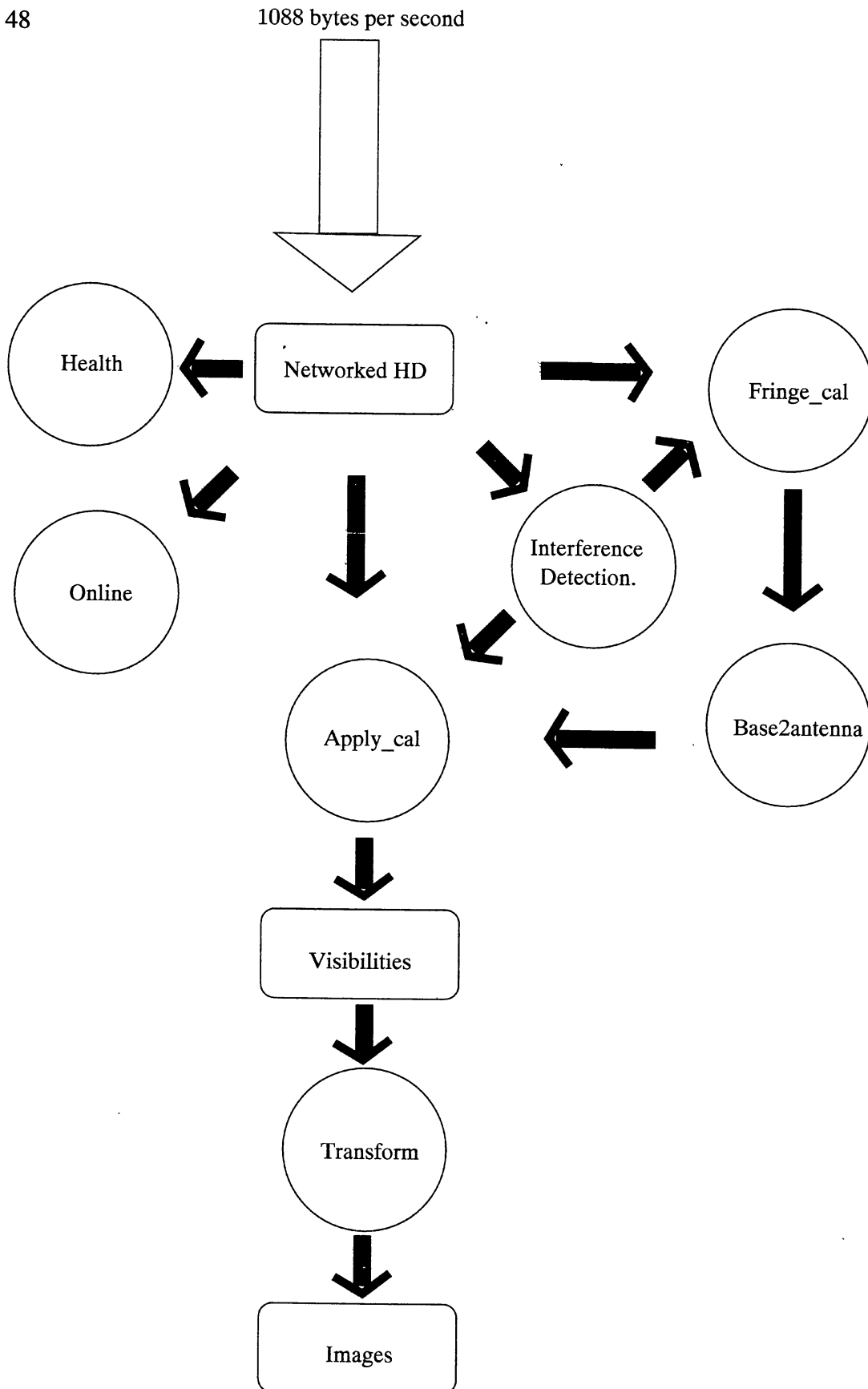


Figure 8. Flow diagram of data acquisition, checking and processing.

interference points which are rejected in subsequent processing. The details of interference detection and their statistics at MRT are given by Sachdev (1998). The programs **Fringe-cal** and the **Transform** are very special to the MRT due to its non-coplanarity and configuration. The following sections describe the concepts behind these programs. Further details of these programs are given by Dodson (1997).

3.1 Antenna calibration

The complex gains of the 47 antennas⁴ are estimated using the measured visibilities for the calibrators MRC0915-119, MRC1932-464 and MRC2211-173. The sensitivity per baseline is 30 Jy for an integration time of one second. In the 1.8° E-W-group primary beam, the above sources are observable in each baseline with a minimum SNR of about 35.

One of the advantages of observing in the southern sky is that the two very strong sources, Cyg A and Cas A are not in the primary beam, where they would dominate the system temperature causing dynamic range problems.

The disadvantage, however, is a paucity of strong sources required for reliable calibration. This leaves us with a situation where we can calibrate the array only a few times in a sidereal day.

The visibility phase is estimated by fringe fitting (hence the name **Fringe-cal** for the program), where we assume that the sky in the primary beam is dominated by the calibrator. The instrumental phase estimated is simply the difference between the phase of the observed visibilities and the expected geometric phase due to the point source calibrator (or calibrators). The instrumental gain is estimated by measuring the relative amplitudes of the fringes on different baselines. The instrumental phase is insensitive to short term interference and also to fringes due to other sources in the sky, provided their fringe rate is significantly different from that of the calibrator for a given baseline. Therefore, at short baselines where there are less than two fringes in the E-W-group beam due to the calibrator, satisfactory calibration cannot be obtained.

This is tackled at MRT by calculating 47 antenna complex gains from the measured 512 complex visibilities (**base2antenna** in Fig. 8). At this stage we reject those baselines with short E-W components (essentially the first four E and W group outputs with the S array) and input the closure information obtained from multiplying the E16 group with the E-W and S array group outputs.

We have found that the day-to-day RMS variation per baseline in phase is about $\pm 7^\circ$ and the RMS amplitude variation is about ± 0.1 dB. The RMS variation of phase from one calibrator to another (e.g. MRC1932-464 and MRC0915-119) is about $\pm 10^\circ$.

Transforming visibilities to brightness: The u, v coverage of MRT (E-W and the N-S components of the baseline) can be thought of as a pleated sheet, extended in both u and v , with discrete steps in w (height) as we move from one E-W group to the one at a different height. As we are imaging a very large field of view (60° field), the approximate coplanar approach, wherein the phase term due to the heights is assumed to be a constant over the synthesized field of view, is invalid. Thus for MRT, a 3D imaging method is required.

⁴ Since the last group of the eastern array (E16) is used as an input to the correlator in the place of the sixteenth trolley.

Here we transform the visibilities using a Fast Fourier Transform (FFT) along the regularly sampled v axis, apply a Direct Fourier Transform (DFT) along the w axis and finally sum along u to obtain the image on the meridian. A DFT on w is required as the sampling is not uniform. The direct transform corrects every term along the zenith angle on the meridian for the group heights. This is equivalent to phasing the groups to a common (and artificial) 2D plane.

The N-S array slopes downwards at about $1/2^\circ$ to the horizontal till about 655 m, and then slopes upwards at about 1° to the horizontal. These two parts are treated separately while transforming. The slope is taken into account by assuming an instrumental zenith appropriate to the slope of the track. While combining contributions from different parts we introduce necessary corrections to get an instrumental zenith equal to the latitude of the place.

The sampling of the visibilities on the E-W grid is at intervals equal to the size of each E-W group. This gives a grating response which falls on the nulls of the primary beam while synthesizing on the meridian. For synthesis away from the meridian, one of the grating response starts moving into the primary beam. This leads to the synthesized beam being a function of the hour angle. Hence to simplify matters, imaging is presently done on the meridian only.

3.2 Initial results

One set of observations for all the 880 trolley positions have been made. Because of the presence of the Sun and the day time interference, the data set is not complete for the full RA coverage. A second set of data is being taken to cover the complete accessible sky. The observations are expected to be completed in 1998.

For imaging on the meridian with all the 880 m baselines, the point source sensitivity will be about 200 mJy (3σ) at the peak of primary beam with 8 seconds of integration and a bandwidth of 1 MHz. The confusion limit for an instrument of resolution $4' \times 4'$ is about 13 mJy. For extended structures the surface brightness sensitivity is about 5 mJy/arcmin².

While the data collection is underway, we are processing the available data to make a low resolution survey of the sky. Observations up to a baseline of 178 m do not require the recirculator. We have taken this as a natural cutoff and have made low resolution images using only 8 central groups of the E-W arm and trolley positions up to 178 m in the N-S. A part of the survey covering the RA range 18:00 to 24:00 hrs and 00:00 to 05:00 hrs has been completed. In Fig. 9, we present a deconvolved image of the region extending from 0100 hrs to 0200 hrs in RA and -62° and -42° in declination, with a resolution of $13' \times 18'$ sec ($\delta + 20.14^\circ$). MRC1932-464, with a flux density of 81 Jy/beam was used as the primary flux calibrator (Golap 1998). The expected RMS noise due to confusion and system temperature (with an integration time of 19 seconds) is expected to be 0.7 Jy. The noise seen on the map is 0.8 Jy and is very close to the expected value. A list of 25 point sources in the flux density range 3 to 20 Jy seen in the map is given in Table 3. Many of them have been identified with Molongolo and Culgoora sources. A comparison of the flux densities of the sources common to MRT and Culgoora lists showed that the error on the estimated flux densities is less than 25%. The vertical line at RA 01h55m and the feature at RA 01h02m, Dec -59° are due to man-made interference. The image has been

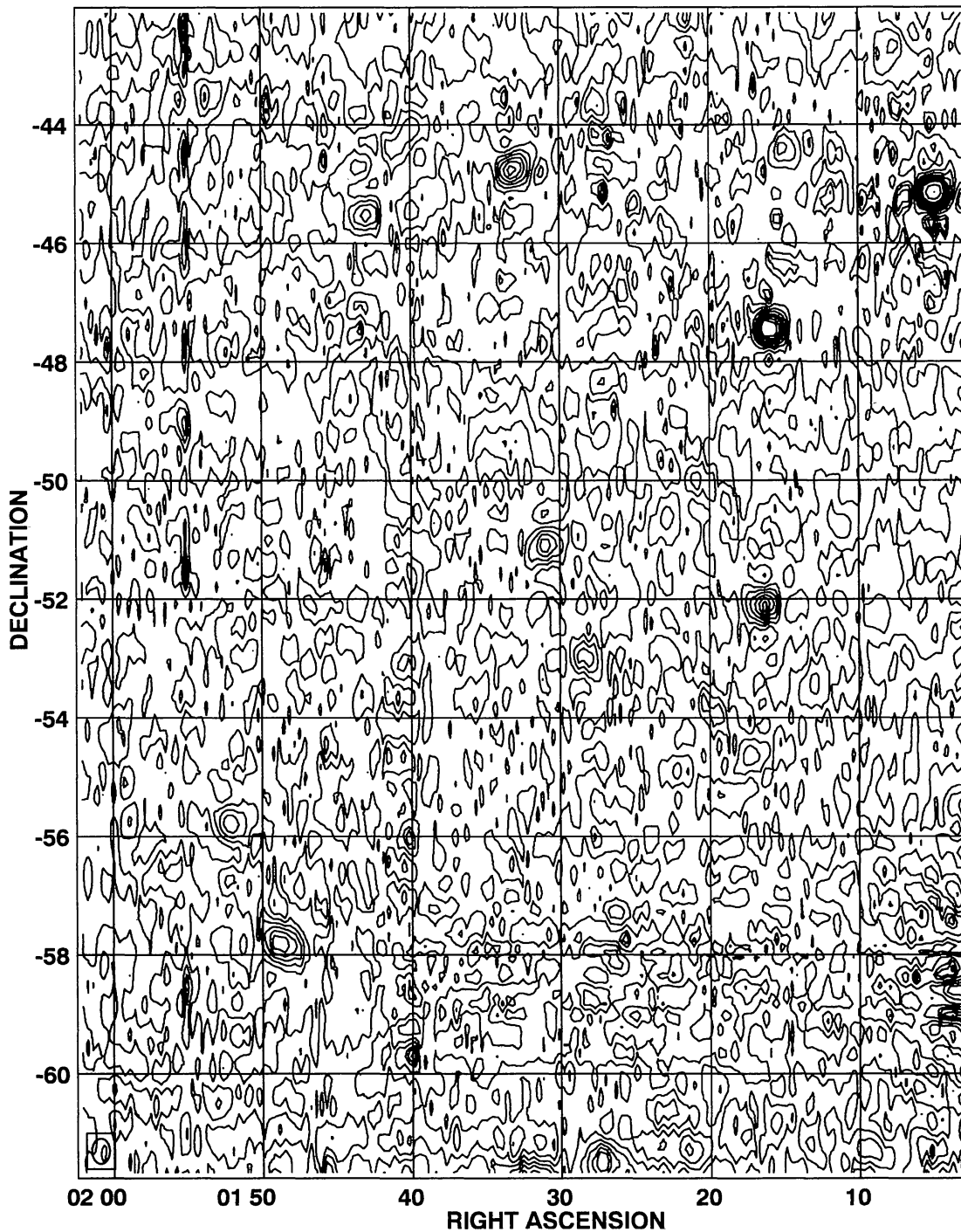


Figure 9. Map at $13' \times 18'$ resolution of a typical region away from the galactic plane. The contours are -9 -7 -5 -4 -3 -2 -1 1 2 3 4 5 7 9 11 13 15 17 19 23 27 31 35 45 55 65 Jy/beam. The coordinates are for epoch 2000.

presented here for completeness and to illustrate the success of the methods developed for imaging. A separate paper on the low resolution survey will discuss more details such as the deconvolution procedure for a non-coplanar array like MRT, method used for detecting point sources, measurement of the zero spacing and the flux density scale used.

Table 3. Table giving point sources detected in the image of Fig. 9. The coordinates are for epoch 2000.

RA	Declination	Flux density (Jy)
01:05:26	−45:04	20
01:07:51	−42:40	6
01:08:53	−52:20	4
01:13:23	−53:23	3
01:15:29	−44:23	4
01:16:30	−47:24	15
01:16:51	−52:08	10
01:17:53	−54:39	3
01:18:54	−62:42	7
01:19:58	−53:48	3
01:21:22	−49:59	4
01:23:27	−50:42	5
01:27:15	−51:37	4
01:27:34	−61:32	9
01:27:57	−48:17	4
01:28:38	−52:58	5
01:31:24	−51:06	8
01:31:44	−61:53	12
01:33:30	−44:40	10
01:43:13	−45:27	6
01:49:04	−57:47	8
01:52:33	−55:49	4
01:54:19	−43:32	5
01:55:41	−49:05	4
01:59:30	−47:30	5

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