

High-resolution, high-sensitivity, ground-based solar spectropolarimetry with a new fast imaging polarimeter

I. Prototype characterization

F. A. Iglesias¹, A. Feller¹, K. Nagaraju^{1,2}, and S. K. Solanki^{1,3}

¹ Max-Planck Institute for Solar System Research, Justus-von-Liebig-Weg 3, 37077 Göttingen, Germany
e-mail: iglesias@mps.mpg.de

² Indian Institute of Astrophysics, Koramangala Second Block, 560034 Bengaluru, India

³ School of Space Research, Kyung Hee University, Yongin, 446-701 Gyeonggi-Do, Korea

Received 24 February 2016 / Accepted 30 March 2016

ABSTRACT

Context. Remote sensing of weak and small-scale solar magnetic fields is of utmost relevance when attempting to respond to a number of important open questions in solar physics. This requires the acquisition of spectropolarimetric data with high spatial resolution ($\sim 10^{-1}$ arcsec) and low noise (10^{-3} to 10^{-5} of the continuum intensity). The main limitations to obtain these measurements from the ground, are the degradation of the image resolution produced by atmospheric seeing and the seeing-induced crosstalk (SIC).

Aims. We introduce the prototype of the Fast Solar Polarimeter (FSP), a new ground-based, high-cadence polarimeter that tackles the above-mentioned limitations by producing data that are optimally suited for the application of post-facto image restoration, and by operating at a modulation frequency of 100 Hz to reduce SIC.

Methods. We describe the instrument in depth, including the fast pnCCD camera employed, the achromatic modulator package, the main calibration steps, the effects of the modulation frequency on the levels of seeing-induced spurious signals, and the effect of the camera properties on the image restoration quality.

Results. The pnCCD camera reaches 400 fps while keeping a high duty cycle (98.6%) and very low noise ($4.94 e^-$ rms). The modulator is optimized to have high (>80%) total polarimetric efficiency in the visible spectral range. This allows FSP to acquire 100 photon-noise-limited, full-Stokes measurements per second. We found that the seeing induced signals that are present in narrow-band, non-modulated, quiet-sun measurements are (a) lower than the noise (7×10^{-5}) after integrating 7.66 min, (b) lower than the noise (2.3×10^{-4}) after integrating 1.16 min and (c) slightly above the noise (4×10^{-3}) after restoring case (b) by means of a multi-object multi-frame blind deconvolution. In addition, we demonstrate that by using only narrow-band images (with low S/N of 13.9) of an active region, we can obtain one complete set of high-quality restored measurements about every 2 s.

Key words. instrumentation: polarimeters – techniques: polarimetric – techniques: image processing – Sun: magnetic fields

1. Introduction

Answering many of the currently open questions in solar physics requires full-Stokes, imaging spectropolarimetric measurements with high spatial (sub-arcsec), spectral (< 100 mÅ) and temporal (< 10 s) resolution along with low noise (10^{-3} to 10^{-5} of the continuum intensity). These types of science questions concern, for example, the total amount of magnetic flux on the Sun (e.g. Trujillo Bueno et al. 2004), the existence and significance of a potential small-scale turbulent dynamo (e.g. Vögler & Schüssler 2007; Buehler et al. 2013; Lites et al. 2014), the structuring and dynamics of the chromospheric magnetic field (e.g. Solanki et al. 2003; Harvey 2009), the contribution of weak internetwork fields to solar irradiance (e.g. Solanki et al. 2013), or magnetic processes related to energy transfer through the solar atmosphere. The interaction of the magnetic field and turbulent convection at the base of the solar atmosphere, structures the field on scales spanning many orders of magnitude in size. There is no doubt today that the plasma processes and magnetic field topologies on scales below 100 km determine the energetics of the higher layers in the solar atmosphere, through the chromosphere and transition region, up to the corona. To acquire a better quantitative understanding of these processes, the small-scale magnetic

field has to be accurately measured, which leads to the challenging requirements mentioned above.

Independent of an observatory's location, on the ground or in space, there are intrinsic trade-offs between these requirements, which arise from the fact that these kinds of measurements are photon starved. The latter being true even for telescopes, post-focus instrumentation, and science cameras with the best possible photon-efficiency (i.e. highest sensitivity and duty cycle).

The main hurdle to obtaining diffraction-limited spatial resolution measurements in visible wavelengths from the ground, are the optical aberrations introduced by the always present Earth atmosphere (seeing). It is mainly the developments during the last 20 yr in terms of adaptive optics (AO) systems (reviewed in detail by Rimmele & Marino 2011) and post-facto image restoration, that have allowed ground-based solar imaging to overcome seeing limitations and reach, although not with the same reliability, a quality comparable to space- and balloon-based observations (Kosugi et al. 2007; Barthol et al. 2011). The two aforementioned techniques, being complementary to each other, have become a standard in most of the state-of-the-art solar facilities and have partially motivated the construction of the next generation of large-aperture (≥ 2 m) solar telescopes (see e.g. review by Kleint & Gandorfer 2015).

The most common image restoration methods applied to ground-based solar data, are Speckle Imaging (Keller & von der Lühe 1992) and different variants of Multi-Object, Multi-Frame, Blind Deconvolution (MOMFBD) see Löfdahl (2002) and references therein. Independently of the technique used, the typical decorrelation time-scale of daytime seeing, sets the maximum exposure time of the individual images to some 10 ms (Löfdahl et al. 2007). The low signal-to-noise ratio (S/N) of such short-exposure frames, makes the restoration of narrow-band images more difficult. This situation can be improved by using an extra high-S/N wideband (WB) channel synchronized to the narrowband science channel (e.g. van Noort et al. 2006; Keller & von der Lühe 1992), or by significantly increasing frame rate and duty cycle of the science channel, as will be shown in Sect. 6.

In terms of polarimetric sensitivity, the main cause for the instrument to depart from its ideal photon-noise limited behavior is the detector noise. The latter can be reduced by minimizing dark current, i.e. cooling the sensor, and by adopting low readout-noise electronics (see Sect. 2). In addition, different spurious polarization signals can harm the measurement accuracy. These include the, sometimes combined (e.g. Sanchez Almeida 1994), effects introduced by instrumental polarization, by detector-related artifacts, and by atmospheric seeing or instrument jitter (Lites 1987; von der Luhe 1988).

Instrumental polarization can be partially reduced either by modeling (e.g. Beck et al. 2005) or compensation (Stenflo 1994; Ramelli et al. 2014). However only the latter helps to fight detector-related artifacts that scale with the level of instrumental polarization, e.g. sensor non-linearities or residual uncorrected offsets (Keller 1996). Moreover, some specific detector artifacts can also be reduced by accurate calibration and modeling (see Sect. 4). On the other hand, there are three methods that are mainly used in ground-based polarimetry to cope with seeing-induced crosstalk (SIC), which can severely harm the accuracy of the inferred magnetic field (e.g. Leka & Rangarajan 2001). These are briefly outlined below.

- (1) AO systems are known to reduce SIC (Judge et al. 2004; Casini et al. 2012). However, the residual artifacts can be large owing to uncorrected high-order aberrations, finite bandwidth of the control system and seeing anisoplanatism. For example, when assuming a single-beam polarimeter modulating at 30 Hz, a Fried parameter of 10 cm and an ideal AO system, which compensates for the first 30 Zernike terms (see e.g. Roddier 1999), the numerical simulations of Krishnappa & Feller (2012) give for the main crosstalk component, from Stokes I to Q , U , and V , a value of about 10^{-3} for a 1 s integration time. For the crosstalk components between Q , U , and V the authors found similar values.
- (2) The combination of (slow) temporal and spatial modulation, particularly in the form of dual beam systems (Lites 1987; Collados 1999), which has successfully been applied to eliminate only the main SIC component (e.g. Collados et al. 2007; Scharmer et al. 2008; Socas-Navarro et al. 2006), and gain a factor of $\sqrt{2}$ in S/N. This at the expense of introducing further spurious signals owing to the differences in the beam path of each channel, e.g. flat field differences, differential image aberrations, etc. These types of beam imbalance artifacts are generally of order 10^{-3} , with some stable systems reaching few times 10^{-4} , e.g. Beck & Rezaei (2009) and Lites et al. (2008). One option, not very common in solar polarimetry, for further eliminating beam imbalance in

dual-beam systems is to use the beam exchange technique (Semel et al. 1993; Bianda et al. 1998; Bommier & Molodij 2002). This technique, however, when used on slow dual-beam systems, has its limitations in terms of high-resolution polarimetry, as it involves combining images recorded at different points in time, which might, in turn, result in a significant loss of spatial resolution owing to seeing.

- (3) Modulating much faster than the seeing to practically “freeze” the atmosphere within a modulation cycle period (Lites 1987). With most of the seeing power contained in the 1–100 Hz (Judge et al. 2004) frequency range, a modulation frequency of order 100 Hz is required to drastically suppress SIC (Krishnappa & Feller 2012).

An efficient modulation regime up to 1 kHz is obtainable with electro-optical modulators based on ferro electric liquid crystals (FLCs, Gandorfer 1999). FLCs are commonly used for full-Stokes polarimetry in the visible or near-infrared. These modulators have a proven longterm reliability and high optical quality, provided that thermal and UV radiation effects are taken into account. Moreover, total polarimetric efficiencies (see del Toro Iniesta & Collados 2000, for a definition) above 80% are commonly obtained, usually by means of a modulator design that is optimized for efficiency across a broad spectral working range (Gisler 2005; Tomczyk et al. 2010).

Piezo-elastic modulators (PEMs, Stenflo & Povel 1985) operate at modulation frequencies of order 10 kHz where SIC becomes completely negligible. PEMs can also be used in the UV below about 380 nm, which is difficult to reach with liquid-crystal based modulators. However, a true full-Stokes PEM-based polarimeter with the required frequency stability has proven difficult to achieve in practice, see for example Gandorfer (1999). Further, the extremely high modulation frequency, which is dictated by the acoustic resonance frequencies of the PEMs, rules out any demodulation by means of a synchronized frame readout, and requires other, much faster demodulation techniques (see below).

The detection of kHz-modulated intensity signals has been the bottleneck of fast imaging polarimeters since the late 1980s, basically because of the inability to get large, 2D detectors with the required cadence and low noise¹. An exception being ZIMPOL (e.g. Povel et al. 1990; Gandorfer et al. 2004), the only imaging polarimeter to date that is able to work in the kHz regime. ZIMPOL uses a dedicated CCD detector which allows on-chip accumulation of the different modulation states and therefore decouples the signal detection from the sensor readout (Povel 1995). This solution renders ZIMPOL fast enough to work with PEMs and has successfully been applied to obtain very high sensitivity ($<10^{-5}$) measurements. However ZIMPOL is not the optimal instrument when high spatial resolution is also a goal (see Sect. 2).

In this work, we describe the prototype of a new, FLC-based, high-cadence, imaging polarimeter developed at the Max Planck Institute for Solar System Research in close collaboration with the semiconductor laboratory of the Max Planck Society. The Fast Solar Polarimeter (FSP) operates in the visible part of the spectrum and was conceived to attain the aforementioned challenging parameter domain, which is required

¹ An alternative approach, based on optical demodulation of the intensity signals before detection, eliminates the necessity of fast sensors (Stenflo & Povel 1985). However, this technique has not been applied because it involves losing a factor of three in intensity and other practical drawbacks like the alignment of three detectors, phase-locked modulator and demodulator, and many optical surfaces.

for high-resolution, high-sensitivity solar polarimetry from the ground. This is possible thanks to a novel high-cadence camera that allows FSP to exploit the benefits of short exposure time, low detector noise, high duty cycle, and fast modulation frequency with synchronous frame readout.

With FSP, we will be able to seriously tackle the science questions mentioned above. For example, FSP will be able to measure the magnetic flux in the quiet Sun via deep Zeeman vector measurements in photospheric lines at about 2–3 times higher spatial resolution than in the infrared lines observed with other polarimeters such as the Tenerife Infrared Polarimeter (TIP). Due to its high sensitivity, FSP can also be used to detect spatial variations of the Hanle depolarisation in the Sr I 460.7 nm line. This will test the hypothesis that the Hanle depolarization of the Sr I line is mainly produced in the intergranular lanes (Trujillo Bueno et al. 2004). By carrying out the first spatially resolved observations of this effect it might be possible to settle the validity of the large amounts of magnetic flux (up to 160 G on average in the quiet Sun, Shchukina & Bueno 2011) previously found using this line.

To address the issue of a small-scale turbulent dynamo, FSP is able to follow up on the technique pioneered by Danilovic et al. (2010), i.e. compare the spatial distribution of Zeeman polarisation in high-resolution observations with that in small-scale turbulent dynamo simulations. Sensitive FSP Zeeman data in the photospheric Fe I line pairs at, for example, 525 nm and 630 nm will allow this technique to be extended to also comparing the evolution with time, which will further strongly constrain the origin of the magnetic features.

Energy transfer through the solar atmosphere is another broad topic where we expect significant contributions from FSP. FSP will be able to look for rapid changes in photospheric and chromospheric magnetic fields, so far undetected owing to the lack of appropriate instruments. This is crucial for detecting high frequency waves, as well as the rapid evolution of magnetic fields in flares. Furthermore, FSP will be able to follow all magnetic features within its field of view to determine how effectively field-line braiding takes place. This requires Stokes vector observations at high resolution and low noise to also capture the weak-field features. Otherwise the amount of braiding will be underestimated, since features will be missed. FSP is predestined to provide the required data.

The description of FSP is divided in two practically independent reports. This work, hereafter refereed to as Paper I, presents the instrumental details and Feller et al. (in prep.), hereafter referred to as Paper II, communicates the results of the first measurements obtained with the FSP prototype.

The present paper is structured as follows. Section 2 compares the main properties of the FSP detector, with respect to the ones used in other visible solar polarimeters to emphasize the need for a high-cadence, low noise sensor. Section 3 describes the instrument and presents the main performance characterization results, including those of the modulation package (Sect. 3.2) and the detector (Sect. 3.3). Section 4 details the data acquisition and the main calibration procedures. Section 5 briefly discusses the measured behavior of SIC at different modulation frequencies. Finally, Sect. 6 presents an example of the application of MOMFBD to FSP data.

2. Detectors used in state-of-the-art, solar polarimeters

As explained in Sect. 1, the intensity detector is a critical component for obtaining the desired properties of a high

spatial-resolution, high-sensitivity, ground-based polarimeter. In Table 1, we have compiled nine important, detector-related properties of the following state-of-the-art, full-Stokes, ground-based polarimeters that work in the visible part of the spectrum²: IBIS (Cavallini 2006); VIP (Beck et al. 2010); CRISP (Scharmer et al. 2008; de la Cruz Rodríguez et al. 2015); DLSP on phase II (Sankarasubramanian et al. 2004, 2006); SPINOR (Socas-Navarro et al. 2006); FSP (see Sects. 3 and 4) and ZIMPOL-3 (Ramelli et al. 2010, 2014). In addition, we included the expected properties of the FSP project in its second phase, labeled FSPII (see Sect. 7).

We note that, for some of the polarimeters listed in Table 1, the adopted values of the detector properties may differ from the maximum allowed by the cameras. The latter is the result of design trade-offs that may include the performance of associated post-focus instruments or different detector parameters, e.g. operating the camera at a reduced frame rate to minimize readout noise. In all cases, we listed the values that are reported as being used during typical polarimetric observations. We excluded from Table 1 the optical filling factors, with most of the sensors having almost 100%; and quantum efficiency, because in general it can be modified by means of sensor coatings to guarantee high (>80%) values in the desired portions of the visible spectrum.

In the following list, we briefly discuss all the properties presented in Table 1, along with their effect on imaging polarimetric performance, and compare their values for the different instruments:

1. Pixel area: two aspects are relevant regarding pixel area. Firstly, detectors with non-square pixels have a more complex point spread function (PSF) and sample the image at different spatial frequencies in the two orthogonal dimensions, further complicating the image restoration process. Secondly, the pixel size has to match the required spatial sampling of the (ideally diffraction limited) image. As a consequence, large pixels may require re-imaging setups of impractical dimensions, which can also introduce further optical aberrations. Therefore, small and square pixels are an advantage. From the second column of Table 1, we note that this aspect constitutes a drawback of the ZIMPOL-3 concept, namely the use of a sensor with three out of four rows covered, which implies the attachment of a micro-lenses array in front of the photosensitive matrix to focus the light on the uncovered pixels and obtain near 100% filling factor. The latter leads to non-square pixels with one large dimension of 90 μm , which is almost a factor of six larger than the most common values (<16 μm) in Table 1, and the necessity of a challenging sub-micron alignment process between pixels and micro lenses (Gandorfer et al. 2004). Besides ZIMPOL-3, the next largest pixel areas, which are two to three times larger than the other listed polarimeters, belong to FSP and FSPII. This

² We excluded from this list the SOLIS/VSM (Keller et al. 2003) although it is an interesting example of early high-cadence polarimetry. The initial idea of the project in 1998 was to use a pair of custom made, 1024 \times 1024, Sarnoff VCCD1024H cameras which had, 18 \times 18 μm^2 pixel area, 46 e⁻ rms readout noise and 300 fps frame rate (Keller 1998). However the system was never delivered by the manufacturer, forcing the VSM team to select the Rockwell HyViSI-1024 model that have 256 \times 1024 pixels and a frame rate of 92 fps (Harvey et al. 2004). Owing to variable dark-levels, inter-quadrant crosstalk and others issues in the Rockwell cameras, VSM was updated in 2010 to use a pair of commercial Sarnoff 1M100 cameras (Balasubramaniam & Pevtsov 2011) similar to the ones used in SPINOR, hence the exclusion from our list.

Table 1. Approximate values of the main, detector-related properties for eight state-of-the-art, full-Stokes, ground-based, solar polarimeters working in the visible part of the spectrum.

Polarimeter	Pixel area [μm^2]	Detector area [px ²]	Frame rate [fps]	Ro. noise [e ⁻ rms]	ADC [bits]	Duty cycle [%]	Mod. freq. [Hz]	Dual beam	WB channel
IBIS	6.8 × 6.8	1024 × 1024	2.86	20	12	3 ^a	0.42	Yes	Yes
VIP	16 × 16	512 × 512	29	5.4	16	100	7.25	Yes	Yes
CRISP	16 × 16	1024 × 1024	37	20	12	70	9.25	Yes	Yes
DLSP*	12 × 12	488 × 652	50	50	14	100	12.5	Yes	No
SPINOR*	16 × 16	512 × 1024	100	40	12	100	12.5	Yes	No
FSP	48 × 48	264 × 264	400 ^b	4.9	14	98.6	100	No	No
FSPII ^c	36 × 36	1024 × 1024	400	5	16	95	100	Yes	Optional ^d
ZIMPOL-3 ^{*e}	22.5 × 90	1252 × 144	1.47	6	16	1.5 ^a	≥1000	No	No

Notes. Instruments indicated with * operate in spectrograph mode, while the others are used in filtergraph mode. FSP has been tested in both modes. See the text for details. ^(a) Computed for an exposure time of 10 ms. ^(b) The FSP camera can achieve a frame rate of 1100 fps using a special readout mode. However, we consider the maximum to be 400 fps because it is the value used during the observations presented here and in Paper II. Such value is also the expected maximum frame rate for FSPII. ^(c) Currently under development, we listed the expected values. See Sect. 7. ^(d) The camera noise is less critical in the WB channel due to the higher flux, thus, it can be implemented with e.g. a fast commercial CMOS camera. ^(e) Even though the CCD55-30 sensor used has frame transfer architecture, the ability of reading while integrating is not implemented resulting in an overhead readout time of about 0.66 s.

is mainly due to limitations of the manufacturing process and may limit their usage in some solar facilities.

2. Detector area: the size and aspect ratio of the detector, directly affect the instrument FOV. In general, large (1k × 1k) detectors are required to image most common scientific targets with sub-arcsec sampling, e.g. active regions. In combination with a spectrograph, a large aspect ratio can be an advantage, depending on the required spectral resolution and range. For filtergraph-based systems, smaller aspect ratios are commonly used.

There are a variety of sizes and aspect ratios listed in the third column of Table 1, with FSP having the smallest dimensions. Large detector dimensions, high frame rate and low readout noise constitutes a general trade-off that depends, among others, on the readout strategy and technology. Cameras that use massive parallel readout are more easily scalable to larger sensor areas, while keeping similar cadence and readout noise to those based on serial strategies. This is the case for the pnCCD (see Sect. 3.3) sensors used in FSP and FSPII, with the small 264 × 264 and the, nearly four times larger, 1024 × 1024 versions being able to reach the same high frame rate (400 fps), low readout noise (≤ 5 e⁻ rms), and very similar duty cycles ($\geq 95\%$).

3. Frame rate: a high frame rate is needed to get a short exposure time, which is required to apply image restoration, without sacrificing duty cycle, and a high modulation frequency to reduce SIC levels (see Sect. 1).

As can be appreciated in the fourth column of Table 1, FSP can reach four times (400 fps) the frame rate of the next fastest polarimeter, SPINOR (100 fps). Given the modulation scheme adopted, the latter implies (see Sect. 3.2), that FSP can take a full Stokes measurement in 10 ms, making it optimal for image restoration and for high-cadence polarimetry of strong signals (see Sect. 6 and Paper II).

4. Readout noise: readout noise is usually the dominant noise source in cooled detectors. If photon noise-limited observations are a goal, when working with low photon fluxes, then adopting a low readout noise camera is of paramount importance. The latter is what discourages the use of commercial high-speed (>kHz) cameras to do high-resolution polarimetry.

From the figures given in the fifth column of Table 1, we can see that VIP, ZIMPOL-3, FSP, and FSPII have readout noises that are approximately one order of magnitude smaller than the rest of the listed instruments. However, among these, FSP and FSPII have frame rates that are more than 13 times higher.

5. ADC: the longer the word of the analog to digital converter (ADC), the lower the quantization noise, which is typically included in the readout noise figure of the camera and added during the digitalization of the sensor voltages (Bennett 1948). We note that, a real ADC will introduce an extra amount of noise and distortion, owing to its non-ideal circuitry, which increases with the conversion rate. The latter is specified by the ADC effective number of bits or ENOB (IEE 2001), which is smaller than the ADC word length (reported in the sixth column of Table 1). For high-speed, high-dynamic-range cameras, the trade-off between digitalization noise and sample rate of the ADC may become an issue, e.g. for a sensor with 50 000 e⁻ full well capacity and a 12 ENOB ADC, the digitalization noise is 3.5 e⁻ rms (Hoslt 1998, Sect. 4.2.5).
6. Duty cycle: high-resolution, spectropolarimetric observations of dynamic signals in the Sun are photon starved, therefore making use of every photon that reaches the detector is crucial. The duty cycle, i.e. the product of the frame rate times the exposure time, is then a determinant parameter. Furthermore, the main motivation for using detectors with full frame transfer architecture (see Sect. 3.3) in polarimetry is their almost 100% duty cycle even at short exposure times. The difference between detectors with and without frame transfer architecture is clearly seen in the Col. 7 of Table 1. Sensors that are not full-frame transfer, e.g. IBIS, or that do not make use of the reading-while-integrating capability of this architecture, e.g. ZIMPOL-3, have overhead readout times that make them almost unusable for polarimetry at both high spatial resolution and high sensitivity, owing to the very low resulting duty cycle.
7. Modulation frequency: as mentioned in Sect. 1, the detection of the modulated intensities can be performed in synchronization with the detector readout, or not. The latter case, requires a special sensor and is the corner-stone of

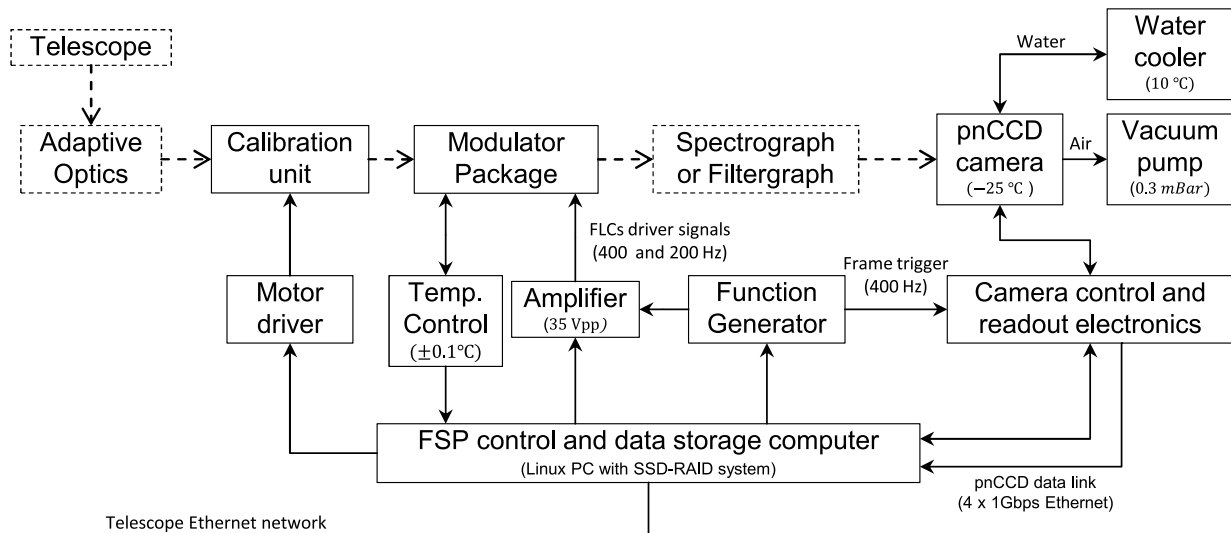


Fig. 1. Components of the FSP prototype. Dashed arrows indicate the beam path. Non-labeled connections using solid arrows represent different control, data, and power lines. Typical operating parameters for some of the blocks are annotated in parentheses. The dashed blocks do not belong to FSP. Two extra communication lines, not shown here, are required for synchronizing the spatial or spectral scanning with FSP acquisition, when a filtergraph or spectrograph is used, respectively.

the ZIMPOL systems. For all the other polarimeters given in Table 1, the modulation frequency is mainly limited by the detector frame rate. Different modulation schemes and technologies require different numbers of intensity measurements N to obtain the full Stokes vector. The modulation frequency is given by the detector frame rate divided by N ; consequently, modulation schemes with low N are to be preferred in this respect³. All the synchronous readout polarimeters listed have $N = 4$ except for IBIS, which has $N = 6$ and SPINOR, which has $N = 8$.

Column 8 of Table 1 details the varied modulation frequency values of the selected polarimeters; with ZIMPOL-3 almost four orders of magnitudes faster than IBIS and ten times faster than the FSPs respectively. We note that the high-frame-rate of the FSP and FSPII cameras and the $N = 4$ modulator, allow these polarimeters to modulate eight times faster than the next fastest polarimeter, which uses synchronous readout (SPINOR), further reducing the SIC levels to very low values (see Sect. 5).

8. Dual beam: as can be appreciated in Col. 9 in Table 1, the dual beam technique is commonly used to reduce SIC in slow polarimeters, since it is not necessary for kHz modulation frequencies. The performance of a fast (>50 Hz) dual beam system strongly depends on the relative values of the residual beam imbalance artifacts and SIC. This point has not yet been addressed in other solar polarimeters and thus motivates the implementation of an optional dual beam setup to be used with FSPII, with the two orthogonal polarization images illuminating different sections of the sensor, i.e. halving the FOV (see Sect. 7).
9. WB channel: as in the case of a dual beam system, acquiring WB images simultaneously to the polarimetric narrow-band data, requires either reducing FOV or the installation of an extra camera with the same frame rate. This is also the case for simultaneous, multi-wavelength observations. In the aforementioned situations, the monetary costs may become

an issue if custom-made, expensive detectors are used in the polarimeter.

In the present section we have compared only cameras that have been tested in solar polarimeters, that are mostly CCD based. However, very-recently-introduced, scientific CMOS cameras (e.g. from manufacturers such as Andor, Fairchild and Hamamatsu), yield very promising specifications in terms of readout noise, frame rate, and quantum efficiency for their application in solar polarimetry (provided that the non-linearities, fixed pattern noise, and any artifact relevant to differential photometry can be controlled).

3. Instrument description and characterization

3.1. System overview

The different components of the FSP prototype are represented in the block diagram of Fig. 1. The various blocks and connections are detailed in the following sections. As can be appreciated, the polarimeter operates in a single-beam configuration. The light (dashed arrows) coming from the output port of the AO system passes through the FLC-based, modulator package – first passing through the instrument calibration unit (ICU) if necessary – to enter the wavelength discriminator system, which can be either a spectrograph or a filtergraph. After this, there is a synchronous detection of the modulated intensities using the pnCCD camera. The signals used to trigger frame acquisition and to drive the modulator are generated by a laboratory function generator. Since the FLCs use a voltage of 35 V_{pp} for a correct switching, a two-channel laboratory amplifier is used. The complete system is controlled by a single Linux-based computer that, in addition, records the data during measurements (see Sect. 4).

3.2. Polarization modulation

The FSP modulator package is similar to the SOLIS/VSM design (Keller et al. 2003). The ideal monochromatic version consists of two FLC, half-wave retarders (FCL1 and FCL2); a quarter-wave, static retarder (SR1); a half-wave, static retarder (SR2);

³ For example, a fast-rotating waveplate polarimeter (e.g. Hanaoka 2012) has $N = 8$, and thus requires twice the frame rate of an FLC-based system, which has $N = 4$, to achieve the same modulation frequency.

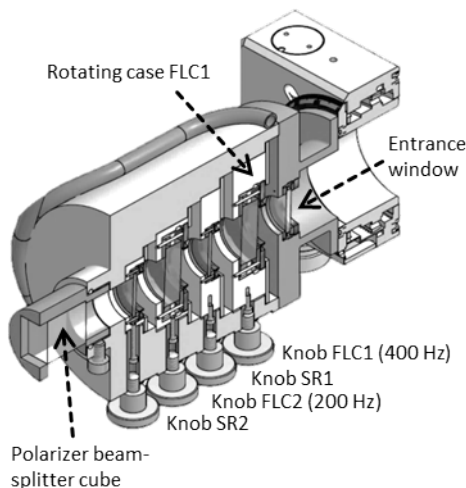


Fig. 2. Cross-section of FSP modulator package. The main components are two ferro-electric liquid crystal (FLC) half-wave retarders (FLC1 and FLC2); a quarter-wave static retarder (SR1); a half-wave, static retarder (SR2); and a polarizer beam splitter cube. The two FLCs and SRs are mounted inside rotating cases to permit the manual modification (by means of the four labeled knobs at the bottom) of their optical axes orientation. The latter allows the optimization of the efficiencies for a given wavelength and/or Stokes parameter in case a specific scientific program requires it.

and a polarizer beam-splitter cube, used as a linear analyzer. The two combined bistable FLCs provide four different modulation states and, thus, the same number of intensity measurements are necessary to obtain the full Stokes vector, i.e. $N = 4$. A cross-section of the modulator is presented in Fig. 2.

The temperature dependence of the FLCs switching angles and retardances (Gisler et al. 2003), demands a thermal stabilization of the modulator assembly to within a fraction of a degree. The large-mass housing and active thermal control, provided by a closed-loop system (see Fig. 1) that is based on a flexible pad heater, ensures a temperature stability of ± 0.1 °C. As a result, a highly stable modulator was obtained, e.g. polarimetric efficiencies that were measured three days apart from each other did not differ by more than 1.5%. The latter significantly reduces the overhead for polarimetric calibration.

The modulator was conceived to minimize the chromatism and unbalance, i.e. the difference between Stokes Q , U , and V of the polarimetric efficiencies in the 400–700 nm wavelength range. Since the manufacturing tolerances of the FLCs can be large, e.g. $\pm 10\%$ for the retardances, we applied the following design procedure to minimize their effect:

- As a first step, two half-wave FLCs were purchased from Boulder Nonlinear Systems and subject to careful characterization. At a given temperature, the dispersion of the retardance, $\phi(\lambda)$, can be expressed in fractions of the wavelength, λ , as (Gisler et al. 2003)

$$\phi(\lambda) = \left(\frac{\lambda_0}{2} - \frac{C_d}{\lambda_0^2} \right) \frac{1}{\lambda} + \frac{C_d}{\lambda^3}, \quad (1)$$

where, λ_0 and C_d are the central wavelength and the dispersion constant, respectively.

- The measured switching angles, λ_0 and C_d for the FLCs (see Table 2), were used as input to obtain a first optimal set of position angles for all the components, and of retardances for the SRs. The optimization procedure, developed

Table 2. Properties of the optical components in FSP modulator.

Property	FLC1	SR1	FLC2	SR2	Unit
Retardance ^a	0.482	0.283	0.453	0.576	λ
Switching ang. ^a	41.5	–	42.6	–	°
C_d ($\times 10^7$)	1.41	0.25	1.41	0.46	nm^3
λ_0	451.9	510.8	434.2	517.7	nm
Position ang. ^b	–71.8	26.7	–41.5	64.8	°

Notes. C_d and λ_0 are model parameters fitted to reproduce the measured dispersion of the retardances at 40 °C (see text). The position angles are the optimum to maximize achromatism and balance of the polarimetric efficiencies in the 400–700 nm wavelength range. ^(a) Measured at 460 nm and 40 °C. ^(b) Defined with respect to the optical axis of the polarizer beam-splitter cube, for the transmitted beam.

by Gisler (2005), simultaneously minimizes, within the 400–700 nm range, the squared differences between the polarimetric efficiencies of a modulator model and the ideal ones for a perfectly balanced and achromatic system, i.e. constant and equal to 1, 0.58, 0.58, 0.58 for Stokes I , Q , U , and V , respectively (del Toro Iniesta & Collados 2000).

- After this, the SRs were ordered from the manufacturer, specifying the desired retardances obtained in the previous step.
- Finally, λ_0 and C_d of the actual SRs that were purchased were measured (see Table 2) and used to run a second optimization with the position angles as the only free parameters to fit.

The final properties of the optical components in the FSP modulator and the optimal position angles found, are summarized in Table 2. The measured spectral response of the polarimetric efficiencies is shown in Fig. 3, along with the results of the numerical modulator model used to perform the optimizations, obtained when the values in Table 2 are input. We also include in Fig. 3 the results of the modulator model after setting all the C_d constants to zero (dotted line). This is done to emphasize the necessity of a detailed modeling of the retarders' dispersions, i.e. Eq. (1), to perform a correct optimization procedure.

3.3. The pnCCD camera

The pnCCD camera was custom-made by PNSensor GmbH together with the MPG semiconductor lab, and is the main component of FSP. It provides simultaneously high frame rate, high duty cycle and low readout noise (see Table 1), allowing the polarimeter to reach a modulation frequency of 100 Hz with synchronous readout. This is accomplished by using a back-illuminated, 3-phase, pnCCD sensor that has a split frame-transfer architecture. Furthermore, all the sensor columns are readout in parallel by means of the 528 amplifiers (integrated in four readout chips) which use eight-fold, multi-correlated double sampling to keep the readout noise low. The measured camera non-linearities are approximately 1% and can be reduced to below 0.1% after calibration. Therefore, polarimetric effects owing to non-linearities (Keller 1996) are in the order of 10^{-5} , given an instrumental polarization of order 1%. The quantum efficiency of the sensor is above 90% in the 500–850 nm wavelength range, and above 60% within 390–1000 nm. Further details on the camera structure and performance are given in Ihle et al. (2012) and Hartmann et al. (2006).

The measured readout plus quantization noise at 400 fps is $4.94 e^-$ rms with a standard deviation over the sensor area of

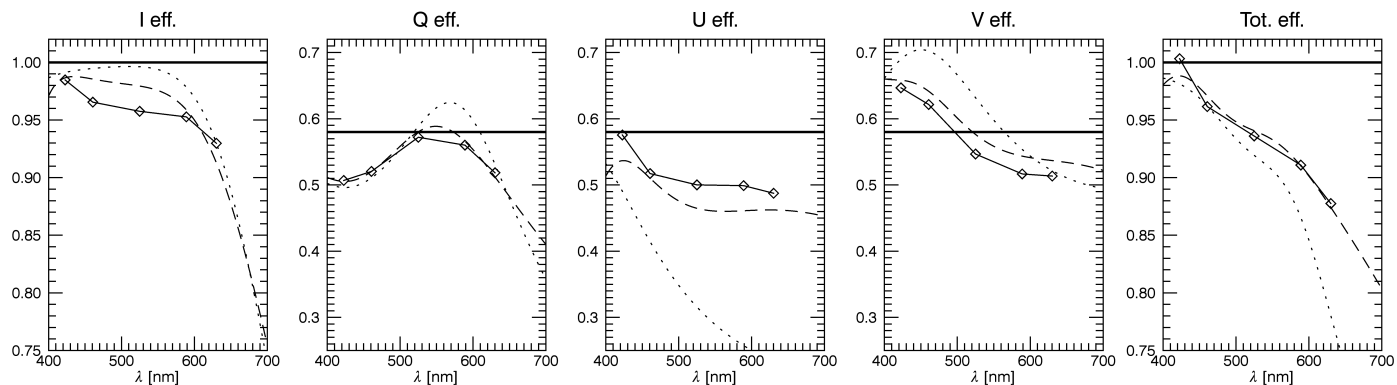


Fig. 3. Measured (diamonds), ideal (thick continuous) and modeled FSP polarimetric efficiencies vs. wavelength. The results of the numerical modulator model are presented for two different configurations, one using the values listed in Table 2 (dashed) and the second using the same input, except that the values of the dispersion coefficients, C_d , were all set to zero (dotted). The large differences of the latter, with respect to the measured values, particularly in Stokes U , emphasize the necessity of a detailed model of the retardances' dispersion to properly reproduce the spectral response of the modulator. The ideal efficiencies are the maximum achievable by a perfectly balanced and achromatic system.

$0.16 e^-$ rms, for a conversion gain of 8.68 digital counts per electron. At the normal camera operating temperature of -25°C , the measured dark noise is $0.31 e^-$ rms. The camera is cooled down using a water-refrigerated (see Fig. 1), thermo-electric Peltier device. As a consequence, part of the camera housing is evacuated (see Fig. 1) down to 0.3 mBar to avoid water vapor condensation on the sensor and electrical contacts.

4. Data acquisition and calibration

FSP can be used in combination with a spectrograph or a filtergraph as post-focus instrument for wavelength discrimination (see Fig. 1). In both cases, the final outcome is a data cube composed of two spatial and one spectral dimensions for each of the four Stokes parameters. Filtergraphs like Fabry-Pérot interferometers, e.g. GREGOR/GFPI (Bello González & Kneer 2008) and VTT/TESOS (Kentischer et al. 1998; Tritschler et al. 2002) are the option of choice in high-spatial resolution observations because the spatial information is recorded strictly simultaneously and, within the isoplanatic patch dimensions of the seeing, degraded by the same PSF. As a consequence, filtergraphs can benefit from post-facto image restoration techniques. On the other hand, spectrographs, which are typically slit-based, can detect a full spectrum simultaneously and be used with long integration times to reach high polarimetric sensitivities, at the expense of reduced spatial resolution. A newly developed technique (M. Van Noort, priv. comm.) promises to allow spectral scans to be restored as well. This would make the use of FSP with spectrographs of considerable interest in the near future.

Independently of the post-focus instrument, the data rate of FSP is large, namely 390 Mbits/s at 400 fps. Consequently four parallel Ethernet connections and a solid-state-disk (SSD) based RAID storage system are used to record the acquired data (see Fig. 1).

4.1. Detector calibration

The following list describes, in order of application, the main data reduction steps used during the calibration of the measured intensities. The steps are independent of the post-focus instrument unless explicitly specified:

1. Offset: the modulated intensity images recorded with the pnCCD camera are corrected for an offset by subtracting

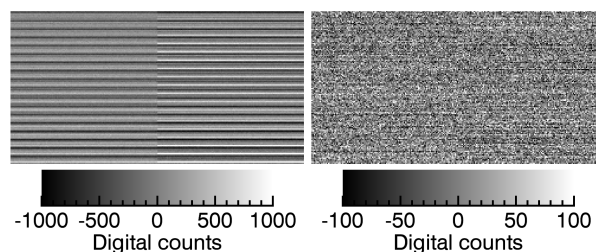


Fig. 4. Offset-corrected dark frame taken with the pnCCD camera, before (left) and after (right) the common mode correction. Only a fraction of the sensor area is shown. The common mode signals are estimated for each semi-row using seven shielded pixels, not shown here, located at the borders of the sensor. The images were taken at 400 fps and -25°C .

a low-noise dark frame, obtained by averaging a long dark series.

2. Common mode: a characteristic issue in parallel readout sensors, such as the pnCCD, is the so-called common mode artifact. This is a variable offset introduced during readout that has, in a given frame, approximately the same value for all the amplifiers belonging to a single readout chip. Therefore, for the pnCCD the artifact appears semi-row wise (see the left image in Fig. 4). Common mode level dominates the image power in dark frames and changes randomly from exposure to exposure, thus it has to be corrected to perform accurate differential imaging. This correction is performed in FSP by subtracting from each semi-row the common mode signal estimated from the average value of seven out of eight shielded pixels, which are located in the borders of the sensor. The pixel closest to the illuminated area is not used owing to the elevated stray light. This reduces the usable area from 264 to 248 columns. Figure 4 illustrates the results of the common mode correction in a dark frame. In addition, Table 3 gives the values of the residual power in dark frames when different numbers of shielded pixels are used to estimate the artifact level.
3. Frame transfer: to keep the duty cycle as high as possible, FSP does not use a shutter. Since the sensor is permanently illuminated during the frame transfer, image smearing can appear (Hoslst 1998). Owing to the high frame rate of the pnCCD camera, the frame transfer takes approximately 1.4% of the exposure time. Therefore, the resulting smearing

Table 3. Measured total noise in dark frames for different numbers of shielded pixels used to correct the common mode artifact.

Num. of pix.	0	1	3	5	7
Total noise	45.32	6.54	5.35	5.07	4.94
Spatial σ	9.65	0.22	0.17	0.17	0.16

Notes. Using zero pixels means that no correction is applied. Both the mean values and the standard deviations of the noise, computed over the sensor area, are given in e⁻ rms. The frames were recorded at 400 fps and -25 °C, with a conversion gain of 8.68 digital counts per electron.

artifact may be of concern for some high-contrast scenes, for this reason, a numerical technique was developed to perform a post-facto correction of the images for the general case of non-constant illumination of the sensor, see [Iglesias et al. \(2015\)](#) for details.

- Flat field: flat-field correction is applied by dividing the four modulated intensities by the same offset, common mode and smearing corrected, normalized flat frame; i.e. polarized flat-field effects are not taken into account. The flat frame is obtained by randomly moving the telescope around the solar disk center. When working with a spectrograph, the spectral lines are numerically removed from the flat frame post-acquisition.

For filtergraph-based observations, after the above-described sensor calibration, the images can be optionally restored using image restoration techniques such as MOMFBD, this is detailed in Sect. 6. The measured intensities are then demodulated using the demodulation matrix obtained from the polarimetric calibration procedure explained in the following section.

4.2. Polarimetric calibration

To perform the polarimetric calibration of FSP, the motorized ICU is introduced in the beam path in front of the modulator package (see Fig. 1). We note this calibration does not include instrumental polarization introduced by the telescope or the AO. However, these are addressed for a specific case in Paper II. The ICU is composed of a high-quality linear polarizer and a super achromatic quarter-wave plate, i.e. retardance is between 77° and 93° in the 400–700 nm wavelength range.

To obtain the polarimeter demodulation matrix, 19 known polarization states are created with the ICU – by rotating its motorized retarder in 10° steps – and measured by FSP. The unknown 4 × 4 elements of the (de-)modulation matrix are then fitted to the measurements. An example of a demodulation matrix obtained at a wavelength of 630.25 nm is presented in Table 4.

To avoid image shifts, which are introduced by the switching of the FLCs (see e.g. [Hanaoka 2006](#)), the modulator package is located close to the scientific focal plane of the telescope. Thus, spatial variations of the demodulation matrix over the FOV are expected. These variations introduce undesired fluctuations of the measured Stokes parameters across the image that scale with the value of the instrumental polarization. Consequently, two options to lessen this effect are (1) the reduction of the instrumental polarization⁴, e.g. given the maximum (three sigma) variation in the FSP demodulation matrix is 2.7×10^{-2} , the instrumental polarization level should be 1×10^{-2} to keep the polarization errors

⁴ Which also has other advantages like the reduction of polarimetric artifacts introduced by sensor non linearities and residual offsets ([Keller 1996](#)).

Table 4. Measured FSP demodulation matrix at a wavelength of 630.25 nm.

34 (0.6)	36 (0.6)	19 (0.5)	11 (0.7)
-23 (0.7)	-54 (0.7)	1 (0.6)	76 (0.8)
67 (0.8)	-45 (0.4)	36 (0.4)	-58 (0.9)
33 (0.5)	27 (0.7)	-82 (0.6)	22 (0.6)

Notes. For each element, both the mean and the standard deviation (in parenthesis) were computed over the sensor FOV and are expressed in %.

near 1×10^{-4} ; and (2) performing the polarimetric calibration and demodulation for each pixel independently.

4.3. Normalization and image accumulation

After demodulation, the polarization levels are normalized with respect to the intensities, i.e. the images of Stokes Q , U , and V are divided by Stokes I . Additionally, many individual Stokes images may be accumulated to increase S/N. The accumulation and normalization operations do not commute and thus the order of their application has to be selected.

For example, in the case of Stokes Q , the two possible options can be expressed for each pixel in the images as $\langle Q/I \rangle$ and $\langle Q \rangle / \langle I \rangle$, where $\langle \rangle$ denotes the average of M images. Considering that both Q and I are noisy quantities with mean values equal to μ_Q and μ_I , respectively, we are interested in the best estimator of the ratio μ_Q/μ_I . Out of the two aforementioned estimators, only $\langle Q \rangle / \langle I \rangle$ is asymptotically unbiased for $M \rightarrow \infty$ (e.g. [van Kempen & van Vliet 2000](#)). This can be recognized after writing the second order approximation to their expected values as follows:

$$\mathbf{E}\left(\left\langle \frac{Q}{I} \right\rangle\right) = \frac{\mu_Q}{\mu_I} + bias \quad \text{and} \quad \mathbf{E}\left(\frac{\langle Q \rangle}{\langle I \rangle}\right) = \frac{\mu_Q}{\mu_I} + \frac{bias}{M}, \quad (2)$$

where $bias = (\kappa\mu_Q - \sigma_{Q,I}^2)/\mu_I^2$, \mathbf{E} represent the expectations value operator, $\sigma_{Q,I}^2$ is the covariance between Q and I , and $\kappa = 1$ Digital count for a Poissonian I . Since μ_I is in the denominator of $bias$, the resulting artifact is non-flat and increases with the frame rate.

The difference between the two estimators is exemplified by Fig. 5, where approximately 5×10^4 Stokes images (corresponding to the measurements of a pore region acquired at -280 mÅ from the line core of Fe I 630.25 nm) were averaged. The images were recorded at 100 Hz modulation frequency using the VTT/TESOS filtergraph and with the FSP modulator deactivated. The latter implies that the Mueller matrix of the modulator is constant in time and thus no real polarization signal is expected⁵, i.e. $\mu_Q = 0$ across the FOV.

5. Spatial resolution and seeing-induced crosstalk

Quantifying the amplitude of SIC in high-cadence, narrow-band, diffraction-limit-sampled, Stokes images, is difficult because the artifact is normally hidden by the photon noise and any possible solar signal. Moreover, if image averaging is used to reduce photon noise, the SIC artifact is smeared and thus its level also reduces (see e.g. [Hanaoka 2004](#), and Sect. 5.1).

⁵ Even though no modulation is taking place, we still use the term modulation frequency to refer to one quarter of the detector frame rate. The same applies to Sect. 5

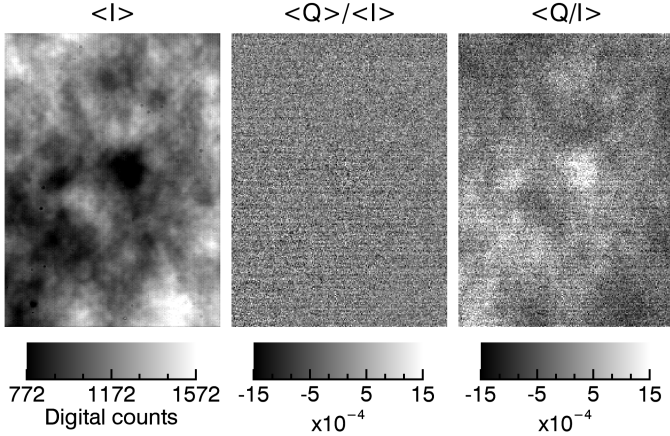


Fig. 5. Average of nearly 5×10^4 (8.3 min) FSP measurements acquired with the modulator package disconnected. Only a fraction of the sensor area is shown. The contrast of the averaged, non-flat-fielded Stokes I image ($\langle I \rangle$) is practically dominated by the etalons cavity errors. The difference of the artifact introduced by the *bias* term of Eq. (2), is evident in the Stokes Q images that were accumulated and then normalized ($\langle Q \rangle / \langle I \rangle$) and vice versa ($\langle Q/I \rangle$).

Let us discriminate the total power in the image of the normalized Stokes parameter $i \in \{Q, U, V\}$, after averaging M measurements as follows:

$$\sigma_{\text{tot}}^2(i, M, f, I_{\text{rms}}) = \frac{\sigma_Y^2}{M\xi_i^2} + \sigma_{\text{art}}^2(i, M, f, I_{\text{rms}}) + \sigma_{\text{sig}}^2(i, M), \quad (3)$$

where we have neglected any fixed pattern noise that cannot be reduced by image averaging. σ_Y is the noise in the measured intensities (normalized to the mean intensity, divided by \sqrt{N} and assumed to be the same in the four modulation states) including the contributions from photon shot noise, dark shot noise, and total readout noise; ξ_i is the polarimetric efficiency of Stokes parameter i ; $\sigma_{\text{art}}(i, M, f, I_{\text{rms}})$ includes the power owing to SIC and other artifacts (see Sect. 5.1) in Stokes parameter i , where we made its dependence explicit with respect to M , to the modulation frequency, f , and to the rms contrast values of the Stokes I images averaged, I_{rms} ; and $\sigma_{\text{sig}}(i, M)$ is the averaged signal power of the actual solar scene.

To be able to reliably detect solar polarimetric signals, one needs to obtain, after averaging, a value of the signal to artifact and noise ratio, $SANR(i, M) = \sigma_{\text{sig}}(i, M) / [\sigma_Y^2 / (M\xi_i^2) + \sigma_{\text{art}}^2(i, M, f, I_{\text{rms}})]^{1/2}$, which is large enough, given a predefined criterion. We note that the above-described approach is of practical relevance only for cases where the portion of the FOV covered by signal and by artifact are similar, so that their spatial rms values are equally representative of their peaks values.

In addition, to preserve spatial resolution, the desired $SANR(i, M)$ should be reached as quickly as possible, i.e. for the smallest M , to avoid the reduction of $\sigma_{\text{sig}}(i, M)$ owing to the spatial smearing produced by solar evolution and seeing. For a given photo-electron flux, the ideal case corresponds to a σ_Y that is photon noise-dominated and $\sigma_{\text{art}}(i, M, f, I_{\text{rms}}) / [\sigma_Y / (\xi_i M)] \ll 1$. As discussed in previous sections, the former is achieved for narrow band images only by adopting a very low noise camera. On the other hand, a low value for the SIC artifact requires a fast modulation frequency.

5.1. Seeing induced crosstalk in quiet Sun images

As an application of the above-expressed ideas, let us consider the case of quiet Sun measurements, which are of particular importance for FSP, given its science goals. One way to achieve $\sigma_{\text{sig}}(i, M) = 0$ in Eq. (3) and still be sensitive to SIC, is to use measurements acquired with the modulator package deactivated (as in Sect. 4.3). In addition, since SIC not only depends on modulation frequency, but also on the instantaneous seeing conditions and contrast of the solar scene (Judge et al. 2004); we use the following procedure to reduce the influence of the latter two.

We start with a set of 45 989 (7.66 min wall time) quiet Sun measurements acquired at 100 Hz with the FSP modulator package deactivated. We then derived two new sets that have modulation frequencies of 20 and 11.1 Hz, by picking one out of five and one out of nine frames, respectively, from the set of measured intensities at 100 Hz. After demodulation, the 20 and 11.1 Hz data sets are equivalent to measurements taken at those frequencies with duty cycles of 19.7% and 11.0%, respectively. The advantages of this method are that practically the same solar scene, seeing conditions, and values of σ_Y / ξ_i are present in the three data sets. The measured values for Stokes I and normalized Stokes Q corresponding to the original and derived measurements are shown in Fig. 6 for different values of M . The qualitative results also apply to Stokes U and V , which are not presented here.

Firstly, we note that SIC is evident in the single measurements, i.e. $M = 1$, of the 20 and 11.1 Hz cases (images c and d respectively), while not being above the noise level for the 100 Hz image (b). Moreover, the artifact tends to have larger amplitudes towards the borders of the FOV because the AO corrections of the wavefront are worse further away from the system's locking point, which is located approximately at the center of the image, owing to seeing anisoplanatism (Rimmele & Marino 2011).

The reduction of the SIC with increasing modulation frequency can be also appreciated in Fig. 7 where the artifact level in the individual Stokes Q measurements,

$$\sigma_{\text{art}}(Q, 1, f, I_{\text{rms}}) = \sqrt{\sigma_{\text{tot}}^2(Q, 1, f, I_{\text{rms}}) - \frac{\sigma_Y^2}{\xi_Q^2}}, \quad (4)$$

is plotted for all the frames, thus including a variety of instantaneous seeing conditions, in the three data sets versus Stokes I rms contrast. To obtain the artifact level, we assume it can only introduce power at spatial frequencies that are below the diffraction limit and that σ_Y is white noise. In which case, σ_Y / ξ_Q can be estimated by the average value of the power located at spatial frequencies beyond the diffraction limit. This estimation of $\sigma_{\text{art}}(Q, 1, f, I_{\text{rms}})$ includes not only SIC contributions, but also any other source of colored noise.

Secondly, from the second row in Fig. 6, we note that SIC has been smeared out to very low levels in the 20 and 11.1 Hz cases (g and h, respectively) owing to image accumulation. This suggests that the SIC artifact is less critical for low-spatial resolution measurements, where, in addition, any small-scale solar signal would also be smeared (compare the contrast of images a and e). If, however, a high spatial resolution is aimed at, reducing the SIC below the photon noise level in the individual measurements is crucial⁶.

⁶ Note that this becomes more difficult when the photon flux reaching the detector increases, e.g. by taking sub-diffraction-limited images with a larger aperture.

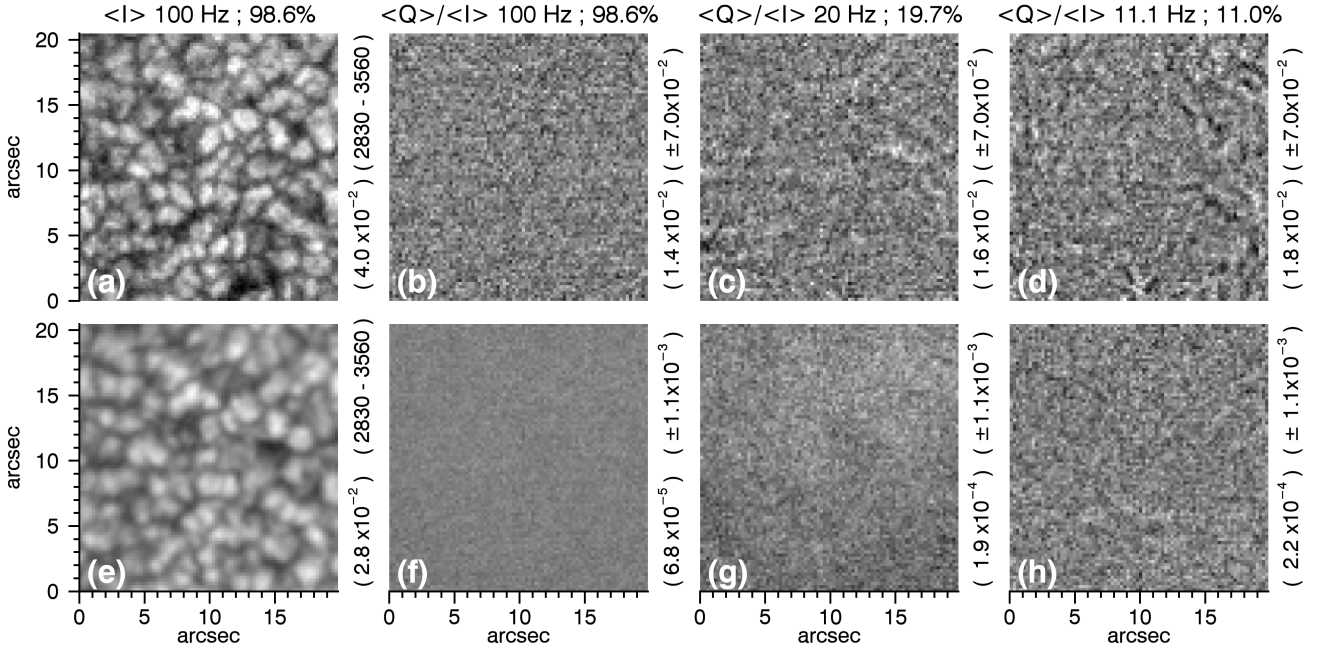


Fig. 6. Seeing-induced crosstalk in a normalized Stokes Q image, for different modulation frequencies, duty cycles (labeled on top of each column) and number of averaged measurements (rows). These quiet Sun images were acquired with FSP modulator package disconnected, i.e. in a situation where no solar polarization signal is expected, using the VTT/TESOS filtergraph tuned to the continuum at -280 mÅ from the line core of Fe I, 630.25 nm. The upper row presents Stokes I **a**) and the normalized Stokes Q **b**), **c**) and **d**) of a single measurement. The lower row shows the results after averaging 45 989 **e**) and **f**), 9 160 **g**) and 5 060 **h**) Stokes images, all covering 7.66 min wall time, approximately. The 20 and 11.1 Hz cases were obtained by eliminating intermediate intensity measurements from the 100 Hz case before demodulation (see the text for details on this procedure). The original plate scale of the images is 0.08 arcsec/pixel in both directions. However they have been binned using a 3×3 pixel² window. The figures in parentheses on the right border of each image denote its grayscale range (upper figure) and standard deviation, or rms contrast for Stokes I , computed across the sensor area (lower figure).

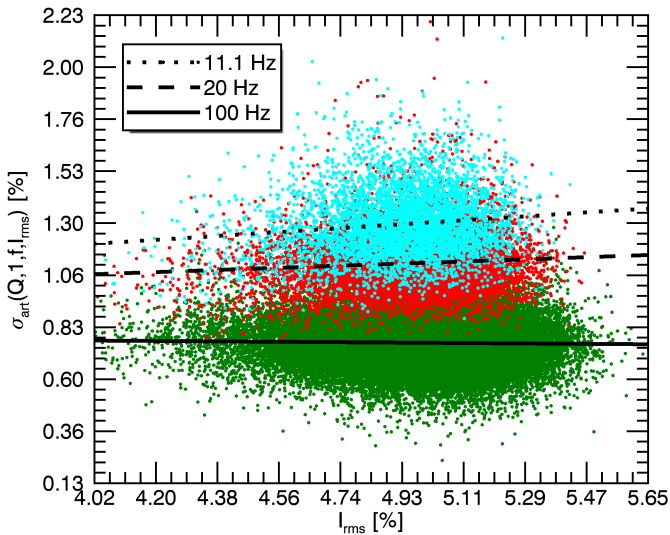


Fig. 7. Measured artifact level in individual Stokes Q measurements (see Eq. (4)) versus Stokes I rms contrast, for the data sets presented in Fig. 6. Each dot gives the value of a single measurement acquired at a modulation frequency of 100 (green), 20 (red) and 11.1 Hz (cyan). The black lines represent linear fits to the 100 (continuous), 20 (dashed) and 11.1 Hz (dotted) data sets.

The latter is also exemplified by Fig. 8, where the resulting Stokes I and normalized Stokes Q , corresponding to the average of 1.16 min wall time of the data sets with 100 and 11.1 Hz modulation frequency, are presented along with their respective

MOMFBD restored versions (see Sect. 6 for details). In the plain averages, both the spatial resolution and the SIC levels are lower because signal (images a and b) and artifact have been smeared out (images e and f). Thus, the benefits of a higher modulation frequency are less prominent. If image restoration is applied to preserve spatial resolution (see images c and d) and any possible small-scale polarimetric signal, then the SIC is also preserved, producing a clear difference between the artifact levels of the cases with 100 and 11.1 Hz modulation frequency. The benefits of a combined high modulation frequency, high cadence, and high duty cycle – to obtain simultaneously high-spatial resolution (through image restoration) and low SIC – are then demonstrated by images c and g.

6. Image restoration

In this section, we present the application of the MOMFBD algorithm, developed by Löfdahl (2002) as an example of post-facto restoration of FSP data. In the multi-frame, blind deconvolution (MFBD) approach, both the object (the constant solar scene) and the degradation functions (the variable seeing plus instrumental PSFs) are simultaneously estimated assuming a noisy, linear shift-invariant degradation model in a maximum-likelihood fashion. To constrain the (initially ill-posed) problem and obtain a more unique solution, multiple short-exposure frames of the same object are recorded, the degradation functions are parametrized, and optical restrictions are added.

The number of frames required for an MFBD is, in principle, small (~ 5), however it depends on the S/N of the individual frames and contrast of the solar scene (Löfdahl et al. 2007). In typical narrow-band, short-exposure measurements, the photon

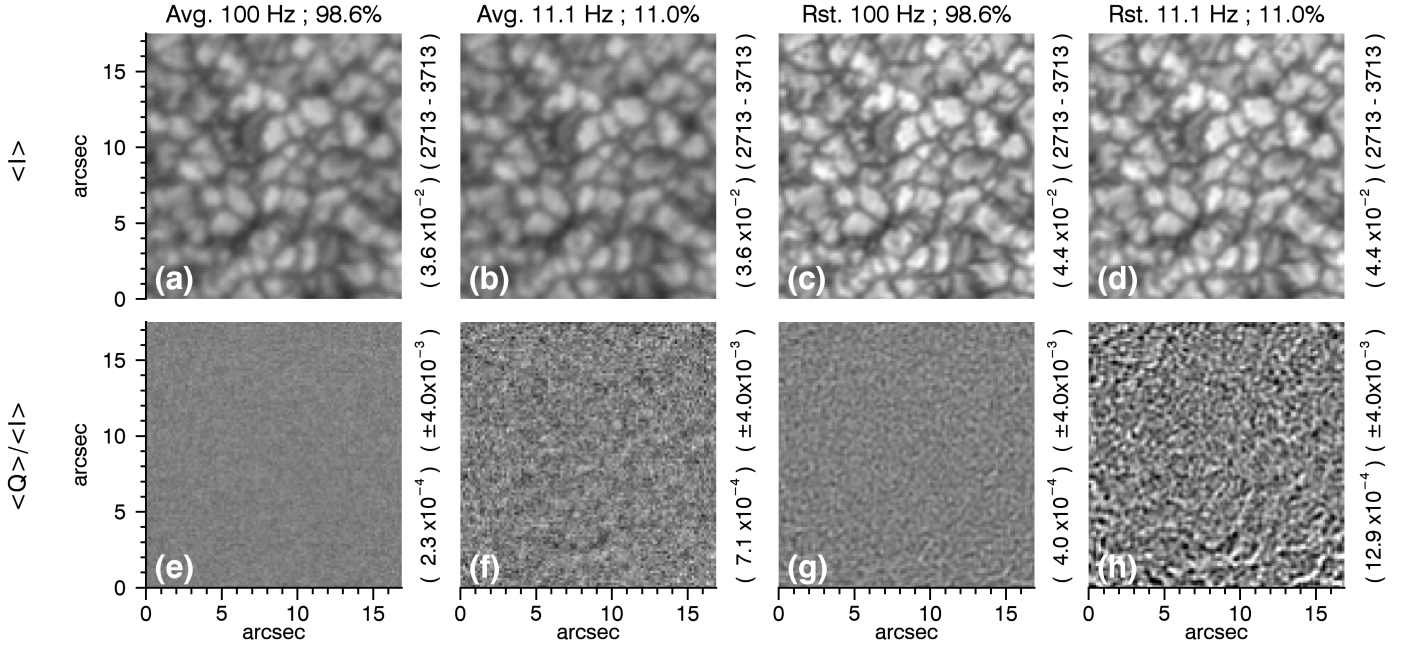


Fig. 8. Effects of the modulation frequency, duty cycle and image restoration (labeled on top of each column) on the SIC level in Stokes Q images. The Stokes I (upper row) and the normalized Stokes Q (bottom row) are shown for the data sets with 100 and 11.1 Hz modulation frequency presented in Fig. 6. Both a plain average, covering 1.16 min wall time **a**), **b**), **e**), and **f**), and the corresponding MOMFBD-restored versions **c**), **d**), **g**), and **h**) are presented. The original plate scale of the images is 0.08 arcsec/pixel in both directions. However they have been binned using a 2×2 pixel² window. The figures in parentheses in the right border of each image denote its grayscale range (upper panels) and standard deviation, or rms contrast for Stokes I , computed across the images (lower panels).

noise is large compared to the signal and thus, if no WB channel is used, an increased number of frames is needed to get a satisfactory restoration. The latter becomes an issue when fast-evolving solar signals are considered because the MFBD algorithm assumes a constant solar scene. In this respect, the duty cycle of the camera is critical to guarantee not only accumulating the maximum number of photons to get a higher S/N, but also a larger number of frames within the evolution time of the targeted solar feature at the desired spatial-resolution.

Since FSP can record the full Stokes vector in 10 ms, i.e. below the typical daytime seeing decorrelation timescale (~ 50 ms), and with 98.6% duty cycle, the resulting polarimetric data can be reduced by (a) restoring the modulated intensities and then performing the demodulation, or by (b) restoring directly the normalized Stokes images. Both approaches have advantages and disadvantages and are suitable for different measurement regimes.

The following is an example of the approach (a) used to obtain a high-cadence (covering 1.92 s) restored image from 192 Stokes measurements acquired with FSP. The restoration was performed using the MOMFBD code implemented by Van Noort et al. (2005). The four groups of modulated intensities, each containing 192 frames, was input as a separate, incomplete object for the algorithm to simultaneously fit independent degradation functions for each modulation state.

The restored Stokes I and normalized V are given in Fig. 9, along with the results of a plain accumulation, to show the improvement not only in spatial resolution (compare d and e) but also in S/N (compare g and h), even when the individual intensity measurements have low (~ 13.9) S/N (see image b). In addition, we include the outcome of the same restoration algorithm, configured in the same way, to a data set with a modulation frequency of 11.1 Hz (44.4 fps) and duty cycle of 11%, which was derived from the original 100 Hz data, as explained in Sect 5.1.

We note that, in this case, the reduced number of frames available for the restoration – only 21 instead of 192 per modulation state – produces a worse restoration (compare e and f), even when the individual intensity measurements have the same exposure time and S/N (compare b and c).

The restorations presented in Figs. 8c–h were also performed as explained above but with one difference, namely that the modulated intensities were treated as simultaneous objects by the MOMFBD code and thus a single degradation function was fitted per modulation cycle. This minimizes possible polarimetric artifacts introduced by the restoration process.

7. Conclusion and prospects

We have introduced the instrumental details for the prototype of a novel, high-cadence solar polarimeter. The FSP is designed to obtain high-resolution, high-sensitivity measurements of the full Stokes vector in visible wavelengths from the ground. To achieve this goal the system uses an FLC-based modulator package optimized to have achromatic and balanced polarimetric efficiencies in the 400–700 nm wavelength range (see Fig. 3). The intensity detection is done in synchronization with the modulator using a small, custom-made pnCCD camera that can record up to 400 fps and has almost 100% duty cycle and total noise below $5 e^-$ rms (see Table 1).

The high modulation frequency (100 Hz) substantially reduces the levels of SIC artifacts per frame to below the noise in typical, narrow-band measurements (see Fig. 6 b). The latter is crucial to achieve simultaneously high spatial resolution (by means of post-facto image restoration) and polarimetric sensitivity (see Fig. 8). Moreover, we found no relevant artifact or systematic effect in quiet-Sun measurements, that were carried out with the modulator package disconnected, that were above 7×10^{-5}

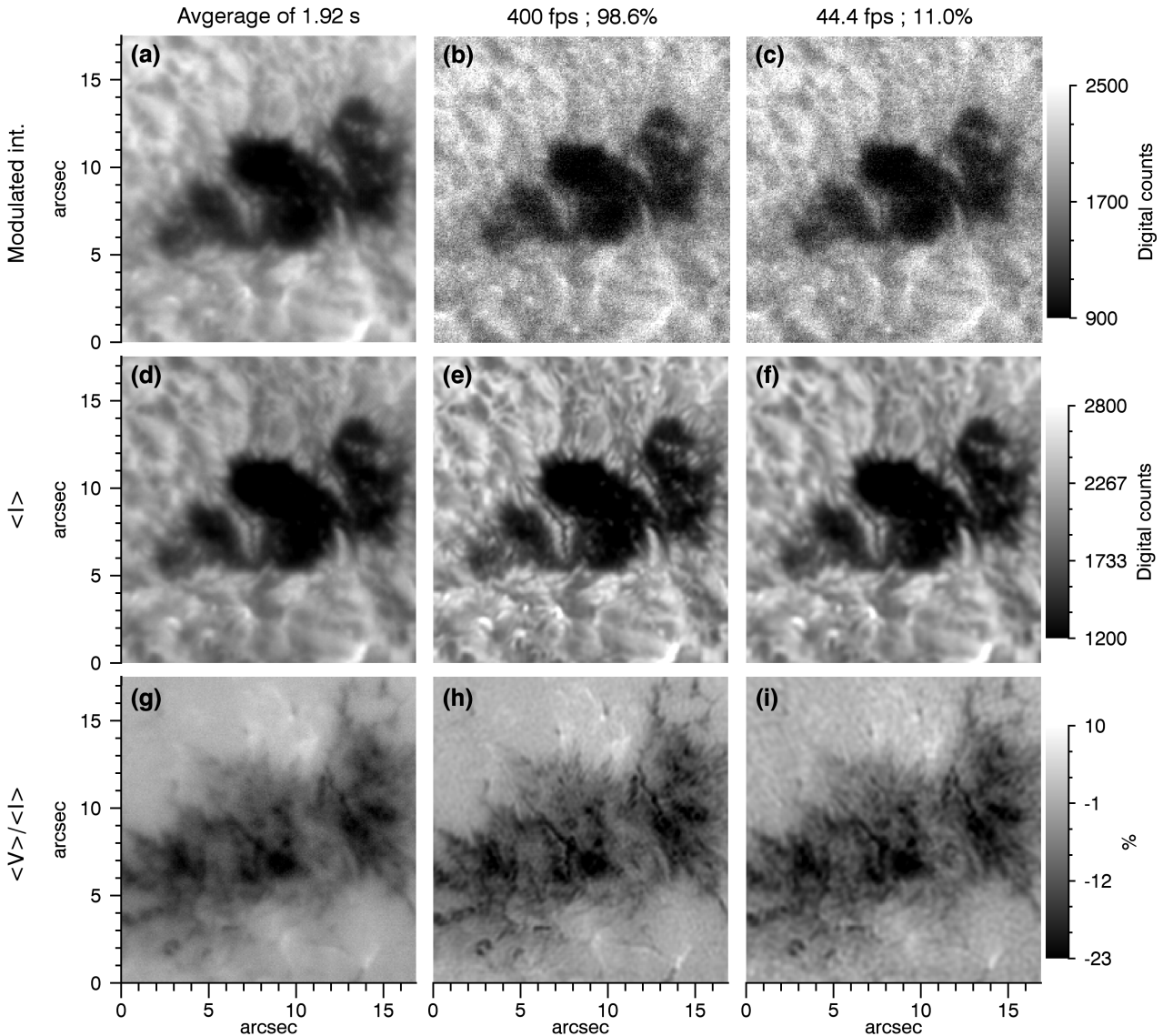


Fig. 9. Effects of the polarimeter cadence and duty cycle when restoring 1.92 s of narrow band data by means of an MOMFBD. The first column presents a simple average of the 192 measurement cycles, including one modulated intensity **a**), Stokes I **d**), and normalized Stokes V **g**). The second column illustrates the results of the MOMFBD restoration performed using all the 4×192 available intensity measurements. Both a single modulated intensity **b**) and two of the resulting Stokes parameters **e**) and **h**) are shown. The third column gives the outcome of the same MOMFBD algorithm, run on the data set that has a reduced frame rate (44.4 fps) and duty cycle (11.1%), including a single modulated intensity **c**) and two of the restored Stokes parameters **f**) and **i**). The data with 44.4 fps frame rate was obtained by eliminating intermediate intensity measurements from the original FSP data set (acquired at 400 fps and a duty cycle of 98.6%) before demodulation. See the text for extra details.

(see Fig. 6f). Similar results were obtained for solar scenes with higher contrasts, e.g. active regions.

Owing to the low noise and high duty cycle of the pnCCD camera, it can produce photon-noise dominated images for all the most common scientific targets tested so far and thus achieve the desired S/N in the shortest time. The high-cadence and duty cycle of FSP are also beneficial when applying post-facto image restoration. The latter was illustrated by restoring a set of 192 measurements using an MOMFBD to show that the larger amount of frames available with FSP improves the restoration quality, even when the S/N of the individual frames is low (see Fig. 9).

The main drawbacks of the FSP prototype are the small sensor area (to be improved in the second phase of the project) and the relatively large pixel size (see Table 1). The latter may result in large optical setups that are not suitable for instruments

with severe space constraints, or that may introduce further image aberrations.

The second phase of the FSP project involves the development of a science-ready instrument. FSPII will use the same modulation package as FSP (described in Sect. 3.2), and a new camera, which is currently under development at the Max Planck Semiconductor Lab. The expected detector specifications are given in Table 1, with the main improvements being the increased sensor area and reduced pixel size. To achieve 400 fps with a four times larger sensor, an improved readout ASIC will be used, namely the VERITAS (Porro et al. 2013). The expected date for the first-light campaign is 2016.

Acknowledgements. The Fast Solar Polarimeter project is funded by the Max Planck Society (MPG) and by the European Commission, grant No. 312495 (SOLARNET). We also would like to thank all the technical contributors not listed as co-authors for their invaluable input to the project. In particular we

would like to thank the staff of PNSensor GmbH and of the MPG semiconductor lab in Munich, for their excellent work on and their support with the pnCCD camera system, as well as the staff of the Kiepenheuer Institut für Sonnenphysik in Freiburg for supporting our FSP observing campaigns at the VTT. This work was partly supported by the BK21 plus program through the National Research Foundation (NRF) funded by the Ministry of Education of Korea. The participation of F. A. Iglesias was funded by the International Max Planck Research School for Solar System Science.

References

- IEE 2001, *IEEE Std 1241-2000*, i
- Balasubramaniam, K. S., & Pevtsov, A. 2011, in *Solar Physics and Space Weather Instrumentation IV*, 8148, 9
- Barthol, P., Gandorfer, A., Solanki, S. K., et al. 2011, *Sol. Phys.*, **268**, 1
- Beck, C., & Rezaei, R. 2009, *A&A*, **502**, 969
- Beck, C., Schlichenmaier, R., Collados, M., Bellot Rubio, L., & Kentischer, T. 2005, *A&A*, **443**, 1047
- Beck, C., Bellot Rubio, L. R., Kentischer, T. J., Tritschler, A., & del Toro Iniesta, J. C. 2010, *A&A*, **520**, A115
- Bello González, N., & Kneer, F. 2008, *A&A*, **480**, 265
- Bennett, W. R. 1948, *Bell System Technical Journal*, **27**, 446
- Bianda, M., Solanki, S. K., & Stenflo, J. O. 1998, *A&A*, **331**, 760
- Bommier, V. & Molodij, G. 2002, *A&A*, **381**, 241
- Buehler, D., Lagg, A., & Solanki, S. K. 2013, *A&A*, **555**, A33
- Casini, R., de Wijn, A. G., & Judge, P. G. 2012, *ApJ*, **757**, 45
- Cavallini, F. 2006, *Sol. Phys.*, **236**, 415
- Collados, M. 1999, in *Third Advances in Solar Physics Euroconference: Magnetic Fields and Oscillation*, eds. B. Schmieder, A. Hofmann, & J. Staude, ASP Conf. Ser., 184, 3
- Collados, M., Lagg, A., Díaz García, J. J., et al. 2007, in *The Physics of Chromospheric Plasmas*, eds. P. Heinzel, I. Dorotovic, & R. Rutten, 368, 611
- Danilovic, S., Schüssler, M., & Solanki, S. K. 2010, *A&A*, **513**, A1
- de la Cruz Rodríguez, J., Löfdahl, M. G., Sutterlin, P., Hillberg, T., & Rouppe van der Voort, L. 2015, *ApJ*, **810**, 11
- del Toro Iniesta, J. C., & Collados, M. 2000, *Appl. Opt.*, **39**, 1637
- Gandorfer, A. M. 1999, *Opt. Eng.*, **38**, 1402
- Gandorfer, A. M., Steiner, H. P. P., Aebbersold, F., et al. 2004, *A&A*, **422**, 703
- Gisler, D. 2005, Ph.D. Thesis, ETH-Zurich
- Gisler, D., Feller, A., & Gandorfer, A. M. 2003, in *Polarimetry in Astronomy*, ed. S. Fineschi, SPIE Conf. Ser., 4843, 45
- Hanaoka, Y. 2004, *Sol. Phys.*, **222**, 265
- Hanaoka, Y. 2006, eds. R. Casini, & B. W. Lites, ASP Conf. Ser., 358, 185
- Hanaoka, Y. 2012, in *SPIE Conf. Ser.*, 8446, 70
- Hartmann, R., Buttler, W., Gorke, H., et al. 2006, *Nucl. Instrum. Meth. Phys. Res. A*, **568**, 118
- Harvey, J. W. 2009, in *Solar Polarization 5: In Honor of Jan Stenflo*, eds. S. V. Berdyugina, K. N. Nagendra, & R. Ramelli, ASP Conf. Ser., 405, 157
- Harvey, J. W., Keller, C., Cole, L., Roy, T., & Jaksha, D. 2004, in *Telescopes and Instrumentation for Solar Astrophysics*, eds. S. Fineschi, & M. A. Gummin, Proc. SPIE, 5171, 258
- Hoslt, G. C. 1998, *CCD Arrays Cameras and Displays* (SPIE Optical Engineering Press)
- Iglesias, F. A., Feller, A., & Nagaraju, K. 2015, *Appl. Opt.*, **54**, 5970
- Ihle, S., Ordavo, I., Bechteler, A., et al. 2012, in *SPIE Conf. Ser.*, 8453, 1
- Judge, P. G., Elmore, D. F., Lites, B. W., Keller, C. U., & Rimmele, T. 2004, *Appl. Opt.*, **43**, 3817
- Keller, C. 1996, in *Solar Polarization*, eds. J. Stenflo, & K. Nagendra (Netherlands: Springer), 243
- Keller, C. U. 1998, in *Advanced Technology Optical/IR Telescopes VI*, 3352, 732
- Keller, C. U., & von der Lühe, O. 1992, *A&A*, **261**, 321
- Keller, C. U., Harvey, J. W., & the SOLIS Team 2003, in *Solar polarization 3*, eds. J. Trujillo-Bueno, & J. Sanchez Almeida, ASP Conf. Ser., 307, 13
- Kentischer, T. J., Schmidt, W., Sigwarth, M., & Uexkuell, M. V. 1998, *A&A*, **340**, 569
- Kleint, L., & Gandorfer, A. 2015, *Space Sci. Rev.*, accepted
- Kosugi, T., Matsuzaki, K., Sakao, T., et al. 2007, *Sol. Phys.*, **243**, 3
- Krishnappa, N., & Feller, A. 2012, *Appl. Opt.*, **51**, 7953
- Leka, K. D., & Rangarajan, K. E. 2001, *Sol. Phys.*, **203**, 239
- Lites, B. W. 1987, *Appl. Opt.*, **26**, 3838
- Lites, B. W., Centeno, R., & McIntosh, S. W. 2014, *PASJ*, **66**, S4
- Lites, B. W., Kubo, M., Socas-Navarro, H., et al. 2008, *ApJ*, **672**, 1237
- Löfdahl, M. G. 2002, in *Image Reconstruction from Incomplete Data II*, eds. J. Bones, M. Fiddy, & R. Millane, 4792, 146
- Löfdahl, M. G., van Noort, M., & Denker, C. 2007, in *Modern Solar Facilities – Advanced Solar Science*, eds. F. Kneer, K. G. Puschmann, & A. D. Wittmann (Göttingen: Universitätsverlag)
- Porro, M., Bianchi, D., et al. 2013, *IEEE Trans. Nucl. Sci.*, **60**, 446
- Povel, H.-P. 1995, *Opt. Eng.*, **34**, 1870
- Povel, H., Aebbersold, H., & Stenflo, J. O. 1990, *Appl. Opt.*, **29**, 1186
- Ramelli, R., Balemi, S., Bianda, M., et al. 2010, in *Ground-based and Airborne Instrumentation for Astronomy III*, eds. I. S. McLean, S. K. Ramsay, & H. Takami, 7735, Proc. SPIE 7735, 1
- Ramelli, R., Gisler, D., Bianda, M., et al. 2014, in *Ground-based and Airborne Instrumentation for Astronomy V*, eds. S. K. Ramsay, I. S. McLean, & H. Takami, Proc. of SPIE, 9147, 3G1
- Rimmele, T. R., & Marino, J. 2011, *Liv. Rev. Sol. Phys.*, **8**
- Roddir, F. 1999, *Adaptive Optics in Astronomy* (Cambridge University Press)
- Sanchez Almeida, J. 1994, *A&A*, **292**, 713
- Sankarasubramanian, K., Lites, B., Gullixson, C., et al. 2004, in *Telescopes and Instrumentation for Solar Astrophysics*, eds. S. Fineschi, & M. A. Gummin, 5171, 207
- Sankarasubramanian, K., Lites, B., Gullixson, C., et al. 2006, in *Solar Polarization 4*, eds. R. Casini, & B. W. Lites, ASP Conf. Ser., 358, 201
- Scharmer, G., Narayan, G., Hillberg, T., et al. 2008, *ApJ*, **689**, L69
- Semel, M., Donati, J.-F., & Rees, D. E. 1993, *A&A*, **278**, 231
- Shchukina, N., & Bueno, J. T. 2011, *ApJ*, **731**, L21
- Socas-Navarro, H., Elmore, D., Pietarila, A., et al. 2006, *Sol. Phys.*, **235**, 55
- Solanki, S. K., Lagg, A., Woch, J., Krupp, N., & Collados, M. 2003, *Nature*, **425**, 692
- Solanki, S. K., Krivova, N. A., & Haigh, J. D. 2013, *ARA&A*, **51**, 311
- Stenflo, J. 1994, in *Solar Magnetic Fields* (Netherlands: Springer), *Astrophys. Space Sci. Lib.*, 189, 312
- Stenflo, J. O., & Povel, H. 1985, *Appl. Opt.*, **24**, 3893
- Tomczyk, S., Casini, R., de Wijn, A. G., & Nelson, P. G. 2010, *Appl. Opt.*, **49**, 3580
- Tritschler, A., Schmidt, W., Langhans, K., & Kentischer, T. 2002, *Sol. Phys.*, **211**, 17
- Trujillo Bueno, J., Shchukina, N., & Asensio Ramos, A. 2004, *Nature*, **430**, 326
- van Kempen, G., & van Vliet, L. 2000, *Cytometry*, **39**, 300
- van Noort, M., Rouppe van der Voort, L., & Löfdahl, M. 2006, in *Solar MHD Theory and Observations: A High Spatial Resolution Perspective*, eds. J. Leibacher, R. F. Stein, & H. Uitenbroek, ASP Conf. Ser., 354, 55
- Van Noort, M., Rouppe van der Voort, L., & Löfdahl, M. 2005, *Sol. Phys.*, **228**, 191
- Vögler, A., & Schüssler, M. 2007, *A&A*, **465**, L43
- von der Luhe, O. 1988, *A&A*, **205**, 354